

Prompt atmospheric neutrino flux

Yu Seon Jeong*

National Institute of Supercomputing and Networking, KISTI, Daejeon 34141, Korea
E-mail: ysjeong@kisti.re.kr

Atri Bhattacharya

Space sciences, Technologies and Astrophysics Research (STAR) Institute, Université de Liège, Bât. B5a, 4000 Liège, Belgium
E-mail: a.bhattacharya@ulg.ac.be

Rikard Enberg

Department of Physics and Astronomy, Uppsala University, Uppsala, Sweden
E-mail: rikard.enberg@physics.uu.se

C. S. Kim

Department of Physics and IPAP, Yonsei University, Seoul 03722, Korea
E-mail: cskim@yonsei.ac.kr

Mary Hall Reno

Department of Physics and Astronomy, University of Iowa, Iowa City, Iowa 52242
E-mail: mary-hall-reno@uiowa.edu

Ina Sarcevic

Department of Astronomy, University of Arizona, 933 N. Cherry Ave., Tucson, AZ 85721
E-mail: ina@physics.arizona.edu

Anna Stasto

Department of Physics, The Pennsylvania State University, University Park, PA 16802
E-mail: ams52@psu.edu

We evaluate the prompt atmospheric neutrino flux including nuclear correction and B hadron contribution in the different frameworks: NLO perturbative QCD and dipole models. The nuclear effect is larger in the prompt neutrino flux than in the total charm production cross section, and it reduces the fluxes by 10% – 30% depending on the model. We also investigate the uncertainty using the QCD scales allowed by the charm cross section data from RHIC and LHC experiments.

*38th International Conference on High Energy Physics
3-10 August 2016
Chicago, USA*

*Speaker.

1. Introduction

Numerous cosmic ray particles are incident into the Earth and collide with the nuclei in the upper atmosphere. From the interactions of cosmic rays with the nuclei, various kinds of hadrons are produced, and some of them subsequently decay generating neutrinos. Neutrinos produced in this process are called atmospheric neutrinos. There are two components in atmospheric neutrinos: conventional neutrinos from the decays of π or K mesons and prompt neutrinos from the decays of heavy hadrons. Although conventional neutrinos are typical atmospheric neutrinos, its flux rapidly decreases with energy due to the energy loss of π or K through interactions before they decay. On the other hand, the heavy hadrons immediately decay, even at high energies, due to their extremely short lifetimes. As a result, the flux of prompt neutrinos dominates at high energy ($E \gtrsim 1$ PeV).

Atmospheric neutrinos are an important background to the astrophysical neutrinos. At present, the IceCube experiment has detected 54 high energy neutrino candidate events between the energies of 20 TeV and 2 PeV, and excluded a pure atmospheric explanation for their origin with high accuracy [1]. The prompt neutrinos can be the primary background to the events at such high energies, hence it is necessary to estimate this flux precisely for the analysis of the experimental measurements.

Among earlier evaluations of the prompt neutrino flux, the prediction by some of us (called ERS) [2] has been used as the benchmark for the analysis of the IceCube measurement. The ERS flux was evaluated based on the so-called dipole model for the heavy quark production. Recently, they re-evaluated the prompt flux in the next-to-leading order perturbative QCD (NLO pQCD) using the updated parton distribution function (PDF), CT10 and more elaborated cosmic ray spectra in [3] (BERSS). The BERSS flux reduces the ERS flux by a factor of 1.5.

In this proceeding paper, we present the improved prompt neutrino fluxes. In particular, we include nuclear corrections and the B hadron contributions with the recent PDF sets and charm fragmentation fractions. We use the experimentally constrained QCD scales to investigate the uncertainty in the prompt neutrino fluxes. Our comprehensive works with more detailed information are presented in ref. [4]. Other recent evaluations appear in [5].

2. Calculation of the atmospheric neutrino flux

The atmospheric neutrino flux can be calculated using the cascade equations, which can be solved by the semi analytic Z -moment method. The cascade equations describe the propagation of particles in the atmosphere in terms of the change of their fluxes. Its general form is given by

$$\frac{d\phi_j}{dX} = -\frac{\phi_j}{\lambda_j} - \frac{\phi_j}{\lambda_j^{(dec)}} + \sum_k S(k \rightarrow j) \quad (2.1)$$

with the interaction (decay) length $\lambda_j^{(dec)}$ and the generation function $S(k \rightarrow j)$. Here, the slant depth X represents the amount of the atmosphere traversed by a particle. The generation function, which depends both on the energy (E) and the slant depth (X), can be rescaled to have only energy dependence. This is called Z -moment, and for particle production, it can be written as

$$Z_{kj}(E) \simeq S(k \rightarrow j) \frac{\lambda_k(E)}{\phi_k(E, X)} = \int_E^\infty dE' \frac{\phi_k(E')}{\phi_k(E)} \frac{\lambda_k(E)}{\lambda_k(E')} \frac{1}{\sigma_{kA}(E')} \frac{d\sigma(kA \rightarrow jY; E', E)}{dE} \quad (2.2)$$

with the assumption, $\phi_k(E, X) \cong \phi_k(E)f(X)$. Then, the solutions of a set of the coupled cascade equations for protons, hadrons and neutrinos can be written in terms of the Z -moments for production of hadrons and those for the decay to neutrinos in the high energy and low energy limits. The final neutrino flux is evaluated by interpolating these two approximate solutions and doing the sum over hadrons. The Z -moment method for evaluating the flux is described in details in ref. [6]. In our evaluation, $D^0, D^+, D_s, \Lambda_c (B^0, B^+, B_s, \Lambda_b)$ were included for the charmed (bottom) hadrons.

3. Heavy quark production cross sections

3.1 Charm/bottom production in perturbative QCD

The essential ingredient for calculating the prompt flux is the production cross sections for heavy quarks, which creates the charmed/bottom hadrons through fragmentation. First we compute the heavy quark production cross section in the perturbative QCD at NLO. Nuclear corrections are incorporated by the recent nuclear PDF, nCTEQ15 [7]. For the factorization (M_F) and renormalization (M_R) scales, we take $(2.1, 1.6)m_q$ for central values, experimentally constrained by the RHIC and LHC data [8].

In Fig. 1 (a), we show the total cross sections per nucleon in pp and $p - \text{nitrogen}$ collisions for the charm and bottom production. The error bands indicate the uncertainty in the cross sections from the scale variation $(1.25, 1.48)m_q$ and $(4.65, 1.71)m_q$ for the lower and upper limits, suggested in ref. [8]. We use this scale range to evaluate the neutrino flux as well. The data points come from the RHIC, LHC and fixed target experiments [9, 10].

The impact of the nuclear correction on the cross section is presented in fig. 2 (a) with the ratio of the cross sections per nucleon for nitrogen and free nucleons. We find that the nuclear effect increases with energy and the cross section is reduced by $\sim 20\%$ (10%) at $E = 10^{10}$ GeV for the charm (bottom) production.

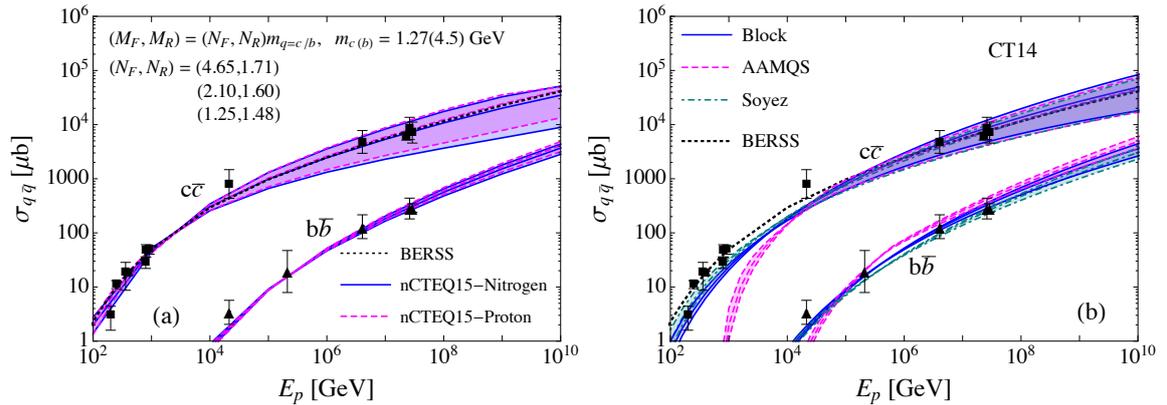


Figure 1: Total cross section for $c\bar{c}$ and $b\bar{b}$ production as a function of the incident proton energy evaluated (a) in the perturbative QCD using the nCTEQ15 PDFs and (b) in the dipole model using the CT14 LO PDF. For comparison the previous evaluation at NLO in pQCD approach is presented as BERSS [3].

3.2 Dipole model

Another possible approach is a dipole model, in which an incoming gluon splits into $q\bar{q}$ pair and interacts with target nucleus as a dipole. The cross section for the heavy quark production in the dipole model can be written as

$$\frac{d\sigma(pp \rightarrow q\bar{q}X)}{dx_F} \simeq \frac{x_1}{x_1 + x_2} g(x_1, M_F) \sigma^{g \rightarrow q\bar{q}X}(x_2, M_R, Q^2 = 0), \quad (3.1)$$

where the partonic cross section $\sigma^{g \rightarrow q\bar{q}X}$ is

$$\sigma^{g \rightarrow q\bar{q}X}(x, M_R, Q^2) = \int dz d^2\vec{r} |\Psi_g^q(z, \vec{r}, M_R, Q^2)|^2 \sigma_d(x, \vec{r}), \quad (3.2)$$

with the wave function squared for the gluon fluctuation, and the dipole cross section (σ_d) for the interaction of the dipole with the target.

We employ the Glauber-Gribov formalism to include the nuclear correction, given by

$$\sigma_d^A(x, r) = \int d^2\vec{b} 2 \left[1 - \exp\left(-\frac{1}{2} A T_A(b) \sigma_d^p(x, r)\right) \right], \quad (3.3)$$

with the impact parameter b . The nuclear profile function $T_A(b) = \int dz \rho_A(z, \vec{b})$ is normalized to unity, $\int d^2\vec{b} T_A(b) = 1$. Here, we use a Gaussian distribution for nuclear density, $\rho_A(z, \vec{b})$.

Fig. 1 (b) shows the total cross sections in pp collisions for the charm and bottom production evaluated using three different dipole cross sections. The Soyez dipole cross section [11] was used for the prediction of the ERS flux in ref. [2]. Other dipoles referred to as 'AAMQS' [12] and 'Block' [13] are more recent and/or improved dipole cross sections. (For the details, see the corresponding references or brief description in ref. [4].) The range of the cross sections are due to the factorization scale in the gluon PDF. Comparing the experimental data, we found that the allowed scale range is $m_c \leq M_F \leq 4m_c$.

Nuclear effects in the dipole models are smaller than in the perturbative QCD, which are about 10 % (2 %) at 10^{10} GeV for the charm (bottom) production. In fig. 2 (a), we present the results only for the Block dipole to avoid cluttering.

4. Prompt atmospheric neutrino fluxes

4.1 Cosmic ray fluxes

In evaluating the prompt neutrino fluxes, we use three cosmic ray fluxes as follows: (a) Broken power law (BPL) – parameterized as $\phi_N(E) = 1.7 (E/\text{GeV})^{-2.7}$ for $E < 5 \cdot 10^6$ GeV and $174 (E/\text{GeV})^{-3}$ for $E > 5 \cdot 10^6$ GeV assuming cosmic rays consist only of protons. (b) H3p – parameterized based on the model for three source components (SN remnants, galactic sources and extragalactic sources) considering the compositions of cosmic rays [14]. Cosmic rays from extragalactic sources are all protons. (c) H3a – similar to H3p, except the mixed composition from the extragalactic sources.

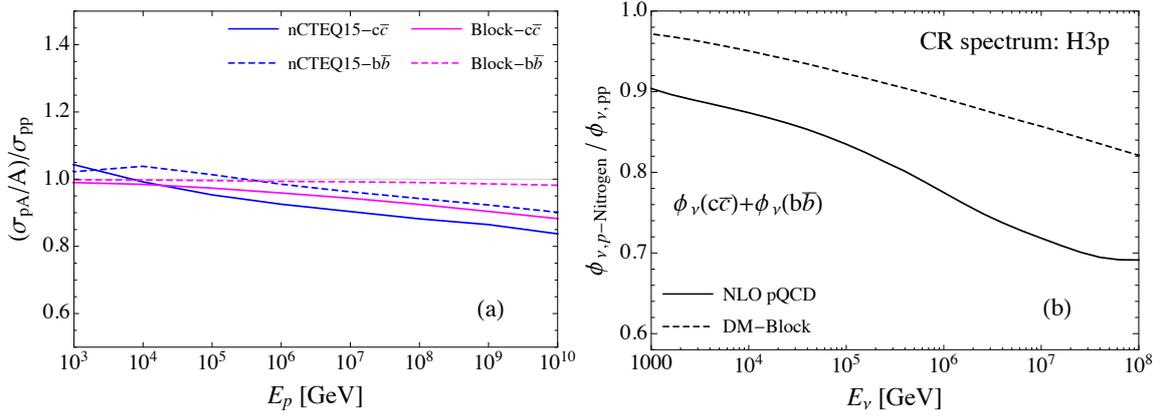


Figure 2: Nuclear effects in the total cross sections for the $q\bar{q}$ production (a) and in the prompt muon neutrino fluxes (b)

4.2 Prompt muon neutrino flux

Fig. 3 shows our resulting neutrino fluxes evaluated with nuclear corrections in the NLO pQCD and the dipole models. For comparison, we present the BERSS and ERS fluxes for the BPL in the pQCD and dipole model results, respectively. For the pQCD case, the error bands are due to the scale variation experimentally constrained by the cross section data as discussed above. Differences from the BERSS flux are due to nuclear corrections, B hadron contributions, the new PDF set and new fragmentation fraction for charm. The combined effect of these results in our new fluxes being $\sim 30\% - 45\%$ lower than the BERSS. The largest impact is from the nuclear corrections, which reduces the flux about $20\% - 30\%$ at $E_\nu = 10^5 - 10^8$ GeV as shown in fig. 2 (b).

For the fluxes from the dipole models, the error bands reflect the uncertainty due to the various dipole cross sections as well as the scale variation. In addition to the updated factors in the pQCD evaluations, the dipole model results have improved with some updated Z moments, Z_{pp} and Z_{DD} , and the charm mass. The nuclear effect in the dipole model evaluations is a $10\% - 20\%$ reduction at $10^5 - 10^8$ GeV. However, as we can see in fig. 3 (b), the fluxes with all the updates are comparable to the ERS fluxes.

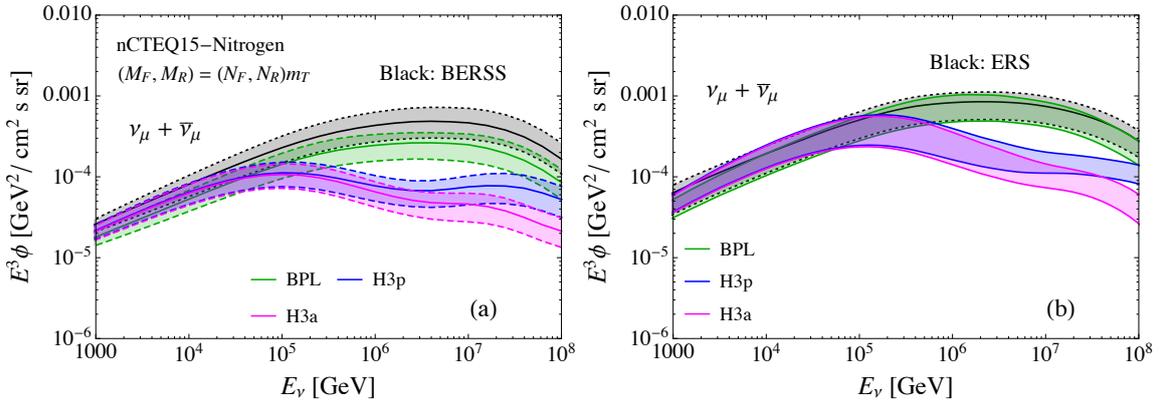


Figure 3: The prompt atmospheric neutrino fluxes evaluated with nuclear correction (a) in the pQCD at NLO and (b) in the dipole models. The BERSS and ERS fluxes are also shown for comparison.

5. Summary and discussion

We have evaluated the nuclear corrected prompt neutrino flux in the different frameworks, NLO pQCD and the dipole models. We also used the QCD scales constrained by the charm cross section data from RHIC and LHC to investigate the uncertainty. The nuclear effect is different according to the models, which is larger in the predictions from NLO pQCD than from the dipole models. The nuclear correction has a more significant impact on the prompt flux, where the forward region is more important, than on the total charm production cross section.

The new prompt neutrino fluxes from NLO pQCD are lower than those from the dipole models, and they do not violate the IceCube limit based on the 3 year observation [1], which excludes most of our dipole model predictions. However, we note that even the fluxes from the dipole models are below the new limit from the 6 year data [15], which was released after our work.

Acknowledgments

This research was supported by the US DOE, the NRF of Korea, the NSC of Poland, the SRC of Sweden and the FNRS of Belgium.

References

- [1] M. G. Aartsen et al. [IceCube Collaboration], arXiv:1510.05223
- [2] R. Enberg, M. H. Reno and I. Sarcevic, Phys. Rev. **D78**, 043005 (2008)
- [3] A. Bhattacharya, R. Enberg, M. H. Reno, I. Sarcevic and A. Stasto JHEP **06**, 110 (2015)
- [4] A. Bhattacharya, R. Enberg, Y. S. Jeong, C. S. Kim, M. H. Reno, I. Sarcevic and A. Stasto, arXiv:1607.00193'
- [5] M. V. Garzelli, S. Moch and G. Sigl, JHEP **1510**, 115 (2015); R. Gauld et al., JHEP **1602**, 130 (2016)
- [6] P. Lipari, Astropart. Phys. **1**, 195-227 (1993)
- [7] K. Kovaric K et al., Phys. Rev. **D93**, 085037 (2016)
- [8] R. E. Nelson, R. Vogt and A. D. Frawley, Phys. Rev. **C87**, 014908 (2013)
- [9] UA1 Collaboration, Phys. Lett. **B256** (1991) 121; LHCb Collaboration, Eur. Phys. J. **C71**, 1645 (2011); ALICE Collaboration, JHEP **07**, 191 (2012); JHEP **11**, 065 (2012); Phys. Lett. **B738**, 97 (2014); ATLAS Collaboration, Nucl. Phys **B907**, 717-763 (2016); PHENIX Collaboration, Phys. Rev. Lett. **97**, 252002 (2006); Phys. Rev. **C91**, 014907 (2015); STAR Collaboration, Phys. Rev. **D86**, 072013 (2012); C. Lourenco and H. K. Wohri, Phys. Rept. **433**, 127 (2006); HERA-B Collaboration, Eur. Phys. J. **C52**, 531 (2007);
- [10] LHCb Collaboration, R. Aaij et al., Nucl. Phys. **B871**, 1-10 (2013); JHEP **03**, 159 (2016)
- [11] G. Soyez, Phys. Lett. **B655**, 32-38 (2007)
- [12] J. L. Albacete et al., Eur. Phys. J. **C71**, 1705 (2011)
- [13] C. Ewerz et al., JHEP **03**, 062 (2011); Y. S. Jeong et al., JHEP **11**, 025 (2014); C. Argüelles et al., Phys. Rev. **D92**, 074040 (2015)
- [14] T. K. Gaisser, Astropart. Phys. **35**, 801 (2012)
- [15] M. G. Aartsen et al. [IceCube Collaboration], arXiv:1607.08006.