

Properties and short-time evolution of nearby galaxies

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We analyze the results of processing of data of observations which had been carried out with the VLBA during 10 last years. All the data have been retrieved from archive of the National Radio Astronomy Observatory (USA NRAO, Socorro, New Mexico). Particularly, we examine data of VLBA observational sessions titled *BK068*, *BL111*, *BL137*, *BL149*, and *BD086*. Objects of our interest are near galaxies with $z < 0.02$. The radio maps of compact structure around the active galactic nuclei (AGN) reconstructed for two such galaxies (*NGC315* and *3C274*) in U frequency band (15 GHz) are presented. Some parameters of these sources are shown in *Table 1*. We have to perform the careful amplitude and phase calibration for all the data. Particularly, a correction of the delay caused by the Earth atmosphere has been made because it is necessary in this frequency range. Bright quasars close to the target sources (for instance, J0136+4751 and 3C279 correspondingly) are used as atmosphere calibrators. Secondly, the Multi Frequency Synthesis (MFS) method (Likhachev, 2004) is used for final maps reconstruction as well as for estimation of spectral index values of core and jet of AGN of both objects. As a result, we can make some conclusions about the milliarcsecond structure of central regions of sources thanks to the using of these two methods of the VLBA data processing. Moreover, we can reveal some changes of this structure in 1999 - 2009, and estimate roughly the velocities and acceleration/deceleration values for brightest components of the nuclei of NGC315 and 3C274. Any polarization phenomena are not taken into account. We present results of processing of data of LL polarization for all the observational sessions

Panoramic Radio Astronomy: Wide-field 1-2 GHz research on galaxy evolution - PRA2009
June 02 - 05 2009
Groningen, the Netherlands

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1. Results of data processing

The software project Astro Space Locator (ASL for Windows) (Chuprikov, 2002) has been used to process all the data. The data processing consists of the following stages :

- Amplitude calibration of all the data using GC and TY tables
- Single band fringe fitting (the primary phase calibration) of all the data. Estimation of the optimum value of solution interval
- Averaging of all the data over each frequency band
- Multi band fringe fitting of the atmosphere calibrator data. Estimation of gain values for every frequency and every time interval
- Application of gains, compensation of atmosphere delay for all the data
- Self-calibration
- Imaging

Figure 1 demonstrates the compact structure of nuclei of 3C274 and NGC315 reconstructed according to the scheme above. Evolution of this structure is obvious for both sources. Table 2 and Table 3 show the changes of coordinates of components. These changes have the following properties :

1. For 3C274

- Image consists of 3 components (A, B, and C)
- There is not any visible proper motion of the brightest component (component A)
- Velocity value of the B component is minimal between February 6, 2003 and February 5, 2007. This value is equal approximately to $1.5 \cdot 10^8$ cm/s , and it increases after February 5, 2007 up to $1.1 \cdot 10^9$ cm/s
- Velocity value of the C component also has a minimum between February 6, 2003 and February 5, 2007. This minimum value is equal to approximately $6.1 \cdot 10^8$ cm/s
Again, this value increases after February 5, 2007 up to $3.6 \cdot 10^9$ cm/s

2. For NGC315

- Image consists of 2 components (A, and B)
- There is not any visible proper motion of the brightest component (component A)
- Velocity of the B component decreases between July 19, 1999 and September 12, 2008 from $6.5 \cdot 10^9$ cm/s down to $1.7 \cdot 10^8$ cm/s

2. Conclusions

We assumed that using the procedure of atmosphere delay calibration can essentially improve the image of source in the 2 centimetre wavelength range. The monitoring of near galaxies radio structure evolution for the period of approximately 10 years is possible thanks to such application of such procedure to VLBA data. Analysis of evolution of structure of two galaxies (3C274, and NGC315) demonstrates absence of superluminal motion of any component in both sources. The

maximum value of component velocity is equal approximately to $6.5 \cdot 10^9$ cm/s (or $0.22 \cdot c$). Thus, we could propose that strong synchrotron self-absorption is significant in both sources. Moreover, spectral index maps show a highly inverted core (spectral index is positive in core of both sources). This could be another indicator of self-absorption.

All the results presented in this paper are preliminary. It is necessary to perform similar monitoring of near galaxies in other frequency ranges to make final conclusions. Procedures and technologies used during the VLBA data processing also could be very useful for processing of data of future Space VLBI mission titled RADIOASTRON.

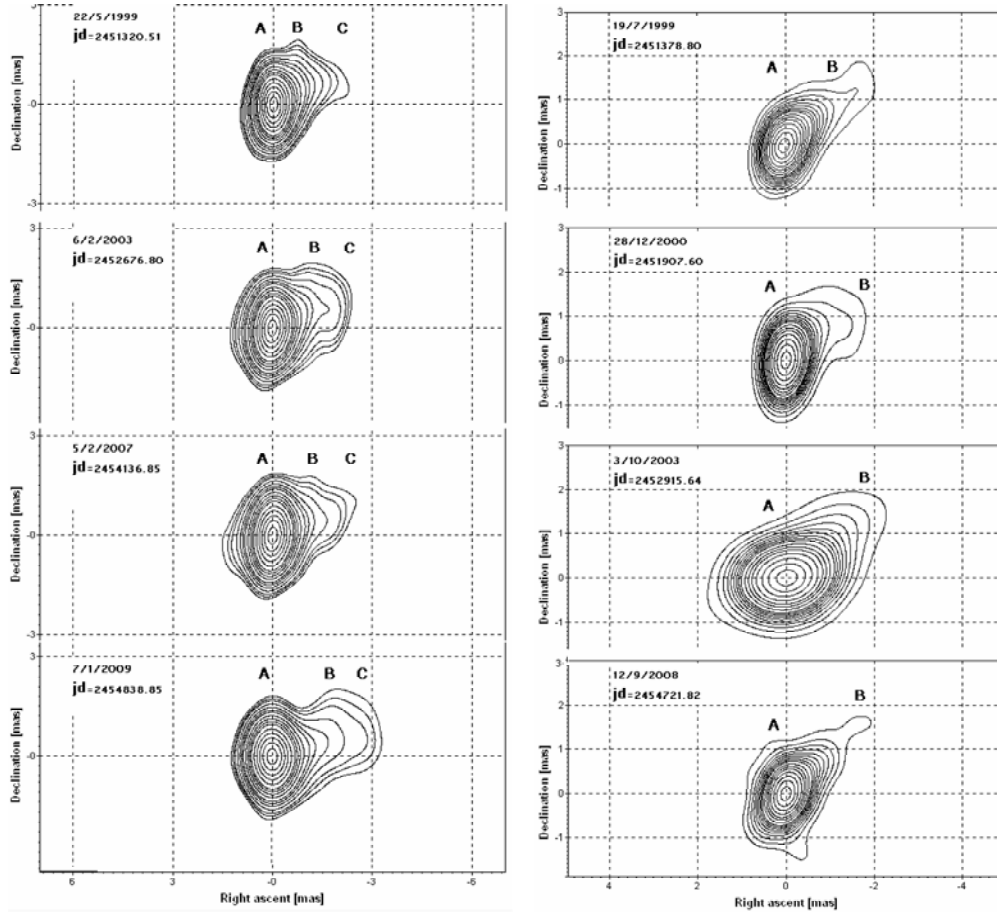


Figure 1: Evolution of the radio structure in 1999 - 2009 for 3C274 (left), and NGC0315 (right)

References

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- [2] E. B. Fomalont 2000, APJSS, 131, 95
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Table 1: Redshift (Fomalont, 2000) and distance for 3C274 and NGC315

IAU Name	Common Name	z	Distance [Mpc]	1 mas corresponds to linear size [cm]
J1230+1223	3C274	0.004	14.7	$2.20 \cdot 10^{17}$
J0057+3021	NGC315	0.017	62.5	$9.35 \cdot 10^{17}$

Table 2: Relative coordinates of A, B, and C components of the central region of 3C274 in 1999 - 2009

Date	JD 2450000 +	VLBA Session	RA of A (mas)	DEC of A (mas)	RA of B (mas)	DEC of B (mas)	RA of C (mas)	DEC of C (mas)
22/05/1999	1320.51	BK068	0	0	-0.870	0.61	-2.32	0.64
06/02/2003	2676.80	BL111	0	0	-1.552	0.59	-2.52	1.01
05/02/2007	4136.85	BL137	0	0	-1.637	0.57	-2.40	0.68
07/01/2009	4838.85	BL149	0	0	-1.900	0.40	-3.35	0.98

Table 3: Relative coordinates of A and B components of the central region of NGC315 in 1999 - 2008

Date	JD 2450000 +	VLBA Session	RA of A (mas)	DEC of A (mas)	RA of B (mas)	DEC of B (mas)
19/07/1999	1378.97	BK068	0	0	-1.618	1.196
28/12/2000	1907.38	BK068	0	0	-1.437	0.930
03/10/2003	2915.65	BD086	0	0	-1.525	1.486
12/09/2008	4721.82	BL149	0	0	-1.500	1.500