

Bursts detected in hard X-rays by the IBIS telescope onboard the INTEGRAL observatory in 2003-2009

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We present results of our search for X-ray bursts in the publically available archival data of the IBIS/ISGRI telescope onboard INTEGRAL. In the January 2003 - April 2009 data we were able to find 833 bursts within the IBIS field of view in the 15-25 keV energy band from 24 known X-ray bursters. Only 243 of the bursts were simultaneously detected by the INTEGRAL/JEM-X telescope in the standard X-ray band. Most of the bursts (586 events) were detected from the well known X-ray burster GX 354-0. Such a large amount of bursts gave us an opportunity to trace the correlations in the burst rate, luminosity, burst duration, fluence and peak flux for this source.

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1. Introduction and data analysis

One of the main evidences of the presence of a neutron star in a binary system is the detection of type I X-ray bursts from it. Type I X-ray bursts originate from thermonuclear flashes taking place on the surface of a neutron star in LMXB (Low Mass X-ray Binary). Depending on the accreted material they last from several seconds to several hours and have a characteristic FRED form with fast rise and exponential decay. Detection of such events may help to identify the nature of new sources and evaluate such parameters of the system as distance, type of accretion material, accretion rate etc. Most of the energy released during type I X-ray bursts is radiated as a black body emission in the standard X-ray energy band. The IBIS telescope onboard INTEGRAL [6] is sensitive to photons with energies over 15 keV. It is a very convenient tool for searching for type I X-ray bursts, because it has a huge field of view (FOV) and dedicates a lot of observing time to study the Galactic Center field where most of the elder stellar population of the Galaxy (including LMXBs) is situated.

The IBIS telescope onboard INTEGRAL consists of two detector layers: ISGRI and PICsIT. In this work we only used data from the ISGRI detector layer in the 15-25 keV energy band. To search for X-ray bursts we first reconstructed detector lightcurves, using all the detected photons irregardless of their incidence direction. These curves were corrected for detector noisy pixels, dead time and turn-offs of one or several detector segments. Then these curves were analyzed for significant excesses in the count rate over its average level. For this purpose we calculated the average count rate for a 500 s window, including each 5 s bin of a given individual observation (SCW - Science Window). Access of the count rate in a single bin over this average value more significant than 3 standard deviations was marked as a burst candidate. Applying this procedure to each 5 s bin of over 60000 SCWs (observations in 2003-2009) we found over 52000 burst candidates. For each of these bins we reconstructed an image of the sky in the FOV of the IBIS/ISGRI telescope again in the 15-25 keV energy range. Comparing statistical significances of the detection of sources in this image with images reconstructed for time bins of the same duration several seconds before and after the current bin we were able to identify 833 bursts with known X-ray bursters (Tabs. 1-4).

2. Results

Similar to previous versions of the catalog [1, 2, 4] over 70% of the bursts were detected from one well known X-ray burster GX 354-0. Our analysis shows that it is not a selection effect, so bursts from this source must be harder than the ones from other sources. One of the possible reasons for this may be Comptonization in the accretion disk corona of thermal photons emitted by the surface of the neutron star during the burst. Five bursters (4U 1724-307, 4U 1702-429, 4U 1608-522, 4U 1636-536 and 4U 1812-12) were responsible for other 22% of the bursts.

Most of the burst sources during the detected bursts were situated outside the FOV of the JEM-X telescope so only 243 bursts were detected in the standard 3-20 keV X-ray band. In principle this may give us a chance to perform spectral analysis for these bursts, but this work is still in progress.

The manual analysis of lightcurves from the IBIS/ISGRI detector revealed several X-ray bursts with statistical significance less than 3 standard deviations, yet well detectable on the corresponding

Table 4: X-ray bursts detected by the IBIS/ISGRI telescope in the 15-25 keV energy range in 2003-2009.

#	Time	Source $\delta t/F^a$	#	Time	Source $\delta t/F^a$	#	Time	Source $\delta t/F^a$	#	Time	Source $\delta t/F^a$
763	Mar 19.82632	GX 354-07/2.3	781	Sep 20.10641	GX 354-05/1.9	799	Oct 2.36093	GX 354-05/2.1	816	Feb 21.91147	GX 354-04/2.1
764	Mar 20.92708	GX 354-08/2.2	782	Sep 20.31375	GX 354-05/1.7	800	Oct 2.47804	GX 354-03/3.0	817	Feb 22.76160	GX 354-07/2.2
765	Mar 22.75968	GX 354-05/2.6	783	Sep 20.71078	GX 354-05/1.9	801	Oct 2.75248	GX 354-04/2.7	818	Feb 22.92381	GX 354-06/1.2
766	Apr 6.13483	GX 354-07/2.1	784	Sep 20.90190	GX 354-06/1.9	802	Oct 5.02817	GX 354-06/1.6	819	Feb 23.09534	GX 354-05/2.6
767	Apr 6.29241	GX 354-05/2.5	785	Sep 21.09396	GX 354-08/2.4	2009			820	Feb 23.26100	GX 354-06/1.9
768	Apr 6.43753	GX 354-03/2.4	786	Sep 21.29137	GX 354-08/2.6	803	Jan 28.63763	4U 1608-522 ¹⁴ /1.7	821	Feb 23.60672	GX 354-06/1.8
769	Apr 6.50782	IJR J17473 ^{f11} /1.3	787	Sep 21.49521	GX 354-08/2.4	804	Jan 29.71948	4U 1608-522 ⁸ /3.0	822	Feb 25.99637	GX 354-07/2.6
770	Apr 6.59962	GX 354-06/2.6	788	Sep 21.71001	GX 354-06/3.5	805	Jan 30.16971	4U 1608-522 ⁶ /2.1	823	Feb 26.36900	GX 354-07/2.8
771	Apr 7.06725	GX 354-05/2.0	789	Sep 21.94872	GX 354-07/1.7	806	Jan 30.57470	4U 1608-522 ⁶ /2.9	824	Mar 4.00832	GX 354-09/1.9
772	Apr 8.57277	GX 354-06/1.3	790	Sep 22.30863	GX 354-08/2.4	807	Jan 31.18564	4U 1608-522 ¹⁰ /2.3	825	Mar 6.64737	GX 354-07/1.7
773	Apr 17.54950	GX 354-03/2.7	791	Sep 23.39289	GX 354-06/1.1	808	Jan 31.59103	4U 1608-522 ⁸ /1.6	826	Mar 6.81059	GX 354-04/1.2
774	Aug 18.08433	GX 354-06/2.3	792	Sep 23.85531	GX 354-05/1.9	809	Feb 1.14700	4U 1608-522 ⁸ /2.2	827	Mar 7.49990	GX 354-06/2.5
775	Sep 10.49925	GX 354-07/2.0	793	Sep 24.21376	GX 354-07/2.3	810	Feb 1.62561	4U 1608-522 ⁷ /1.5	828	Mar 11.40462	4U 1608-522 ¹⁰ /3.2
776	Sep 11.12200	GX 354-05/2.0	794	Sep 24.38043	GX 354-06/1.1	811	Feb 5.06657	4U 1608-522 ¹⁰ /1.8	829	Mar 12.59326	GX 354-07/1.2
777	Sep 18.44472	GX 354-06/2.5	795	Sep 24.59205	GX 354-08/2.5	812	Feb 6.43797	4U 1608-522 ¹⁰ /2.3	830	Mar 19.61958	4U 1608-522 ¹¹ /2.1
778	Sep 18.63818	GX 354-05/1.8	796	Sep 24.77784	GX 354-06/2.3	813	Feb 7.10690	4U 1608-522 ¹¹ /2.3	831	Apr 2.85838	4U 1636-536 ⁶ /1.0
779	Sep 18.83859	GX 354-08/2.9	797	Sep 24.95779	GX 354-06/1.0	814	Feb 7.59847	4U 1608-522 ¹⁰ /2.3	832	Apr 9.55622	GX 354-05/1.0
780	Sep 19.06193	GX 354-03/1.9	798	Sep 25.12663	GX 354-06/3.0	815	Feb 7.91425	4U 1636-536 ⁷ /1.9	833	Apr 18.01957	GX 354-05/0.8

^a - burst duration (in seconds) / pick flux (in Crab); ^b *SAX J1712.6 – 3739*; ^c *IGR J17380 – 3749* [5]
^d *AX J1754.2 – 2754* [3]; ^e *SAX J1810.8 – 2609*; ^f *IGR J17473 – 2721*

sky images. Therefore, although it increases computing time an order of magnitude, we plan to search and identify all X-ray bursts with statistical significance exceeding 2 standard deviations.

Other possible applications of such a large number of detected X-ray bursts include investigation of statistical properties of the phenomena and empirical study of the burst rate - accretion rate dependence. This work is in progress and we are planning to submit a paper with the results by the end of this year. Due to the lack of volume here we list the minimum information on the bursts (Tabs. 1-4). The complete list of burst parameters will appear in the above-mentioned paper.

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References

- [1] I.V. Chelovekov, S.A. Grebenev and R.A. Sunyaev, *Astronomy Letters*, 2006, 33, 456-477
- [2] I.V. Chelovekov, S.A. Grebenev and R.A. Sunyaev, *Proc. of the 6th INTEGRAL Workshop*, 2007, ESA SP-622, 445
- [3] I.V. Chelovekov and S.A. Grebenev, *Astronomy Letters*, 2007, 33, 807-813
- [4] I.V. Chelovekov and S.A. Grebenev, *Proc. of the 7th INTEGRAL Workshop*, 2008, *PoS(Integral08)034*
- [5] I.V. Chelovekov and S.A. Grebenev, *Astronomy Letters*, 2010, 36, 895-903
- [6] C. Winkler, T.J.-L. Courvoisier, G. Di Cocco, et. al., *Astron. Astrophys.* 411, L131