

The Dense Stellar Systems Around Galactic Massive Black Holes

Roberto Capuzzo-Dolcetta*

Dep. of Physics, Sapienza, univ. of Roma, Italy

E-mail: roberto.capuzzodolcetta@uniroma1.it

Manuel Arca-Sedda

Dep. of Physics, Sapienza, univ. of Roma, Italy

E-mail: manuel.arcasedda@uniroma1.it

Mario Spera

Dep. of Physics, Sapienza, univ. of Roma, Italy

E-mail: mario.spera@uniroma1.it

The central regions of galaxies show the presence of massive black holes and/or dense stellar systems. The question about their modes of formation is still under debate. A likely explanation of the formation of the central dense stellar systems in both spiral and elliptical galaxies is based on the orbital decay of massive globular clusters in the central region of galaxies due to kinetic energy dissipation by dynamical friction. Their merging leads to the formation of a nuclear star cluster, like that of the Milky Way, where a massive black hole (Sgr A*) is also present. Actually, high precision N - body simulations (Antonini, Capuzzo-Dolcetta et al. 2012, [1]) show a good fit to the observational characteristics of the Milky Way nuclear cluster, giving further reliability to the cited *migratory* model for the formation of compact systems in the inner galaxy regions.

*Nuclei of Seyfert galaxies and QSOs - Central engine & conditions of star formation,
November 6-8, 2012
Max-Planck-Institut für Radioastronomie (MPIfR), Bonn, Germany*

*Speaker.

1. Introduction

Compact Massive Objects (CMOs) are almost ubiquitously present in the galactic centers, in form of massive or supermassive black holes (MBHs, SMBHs) or in a more *dilute* form, i.e. resolved stellar nuclei or Nuclear Star Clusters (NSCs). There is a direct correlation between compactness of the CMO and the host galaxy luminosity: the brighter the host the denser the CMO, till the limit of BH infinite density (see, e.g., [11]).

An interpretation of the modes of formation of CMOs in galaxies is surely a modern topic, which have not yet received an adequate self-consistent explanation frame. Preliminarily, we must say that a common origin for the CMOs in galaxies, with a second order difference in their compactness and mass as due to properties of the host galaxy, is an appealing view, but faces with the evidence, for example, that some galaxies contain both an NSC and a SMBH [3]. A well-known example is the Milky Way (MW), where a $4 \times 10^6 M_{\odot}$ black hole coexists with an NSC ~ 4 times more massive ($M_{NSC} \approx 1.5 \times 10^7 M_{\odot}$); the same coexistence has been found recently in various other cases. Anyway, in spite of these peculiarities, it is surely intriguing the possibility, raised by Capuzzo–Dolcetta [4] and furtherly developed by him and collaborators, that CMOs have formed by stellar aggregations like massive globular clusters (GCs) brought to the host galaxy center via dynamical braking (dynamical friction). The infalling and orbitally decaying GCs eventually merge in the innermost region of the host galaxy where they form an actual Super Star Cluster, in times and modes that are, partly, governed by the galaxy structure, i.e. geometrical shape, steepness of mass density distribution as well as total mass. If the two phenomena of dynamical friction braking, yielding to a central mass accumulation, and tidal disruption, working in the opposite direction, are the leading mechanisms, the host galaxy luminosity (mass) correlation with the CMO compactness has a qualitative explanation. Actually, fainter galaxies show a higher central density in phase space, thus enhancing the effect of dynamical friction on GCs and favoring their accumulation to grow a NSC, while bright and very bright galaxies borrow, usually, MBHs and SMBHs in their center which may act as destroyers of incoming GCs inhibiting the local NSC formation.

2. The formation of the MW Nuclear Star Cluster

The possibility of cumulation of a significant amount of mass in the innermost regions of galaxies in form of infalling GCs is an old idea [9]. For some time, the dynamical friction effect was considered important just for very massive GCs, until its role, enhanced in triaxial galaxies [8] and in cuspy galaxies [10, 2], was clearly shown to be relevant for a wide range of GC structural parameters and initial conditions by [4] and later papers. This framework of GC infalls provides possible answers on both the origin of local black hole mass feeding and on the modes of formation of the observed dense stellar systems in both dwarf ellipticals and spirals. Actually, these systems have characteristics similar to those of GCs but on a magnified scale. This *migratory* explanation is alternative to that claiming a funneling of molecular gas toward the inner part of galaxies and subsequent local star formation (see for instance [6]). The two hypotheses have been differently tested by various authors; the *dissipationless* hypothesis (GC migratory model) has also been quantitatively well tested by specific N -body simulations, while the *dissipative* (gaseous) origin has not yet been thoroughly investigated.

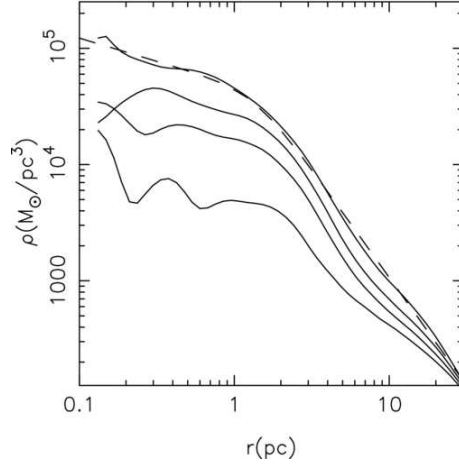


Figure 1: Density profile of the NSC after 3, 6, 9, and 12 infall events. The central density grows with time. The dashed line is the fit to the NSC density profile obtained at the end of the entire simulation using a broken power-law (from [1]).

It seems, so, logical performing a test of the migratory model for the origin of galactic central dense stellar systems in the best known environment: the Milky Way, where, around the $\sim 4 \times 10^6 M_{\odot}$ central black hole there is a NSC of about $10^7 M_{\odot}$ whose photometric structure as well as kinematical characteristics are studied at a level unreachable in external galaxies. As a matter of fact, the MW NSC is, so far, the only one resolved in its stellar content in spite of its high density. As a consequence, in the recent paper [1] tackled the problem of analyzing whether the migratory hypothesis for the formation is working also in the specific context of the MW, where, indeed, the dynamical role of the black hole is likely to be significant.

The observational constraints are:

i) the mass of the NSC ($\sim 1.5 \times 10^7 M_{\odot}$); *ii)* its mass density profile, $\rho(r)$, showing a core of 0.5 pc size in the distribution of the late type stars, and decreasing further out as $\rho \propto r^{-1.8}$ up to $r \simeq 30$ pc. Out of 30 pc there is a large nuclear stellar and molecular disk of about same radius; *iii)* some kinematic data; *iv)* the NSC luminosity function.

We showed how the orbital decay of 12 tidally truncated massive ($\sim 10^6 M_{\odot}$) GCs actually leads, in a sufficiently short time, to the growth and stabilization of a dense stellar cluster around the Sgr A* BH. We followed the infall and merger of the GCs in two sets of initial conditions (corresponding to simultaneous and consecutive infalls) and were able to study the evolution of the density profile as well as the distribution in velocity space. Moreover, we studied also the post merger evolution, extending the N -body simulations after the final merger for a time that corresponds to about 10 Gyr, after scaling to the MW. The absence of a clear formation of a Bahcall-Wolf cusp in the NSC is not surprising because the expected time scale for its growth is the 2-body relaxation time at the black hole influence radius is in the range 20 – 30 Gyr [7].

3. Results

Taking into account the constraints imposed by the observations of the MW NSC, our results

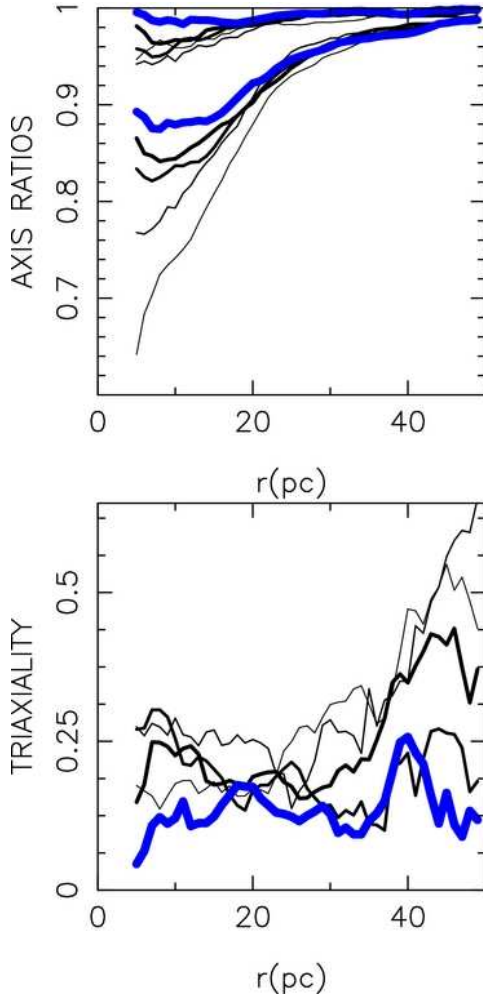


Figure 2: Evolution of the model axis ratios (upper panel) and triaxiality parameter T (lower panel), as functions of radius, in the post-merger phase. Line thickness increases with time; the times plotted span the 0.12 - 0.6 range, in units of relaxation time at the influence radius. Thickest blue lines correspond to the final model (from [1]).

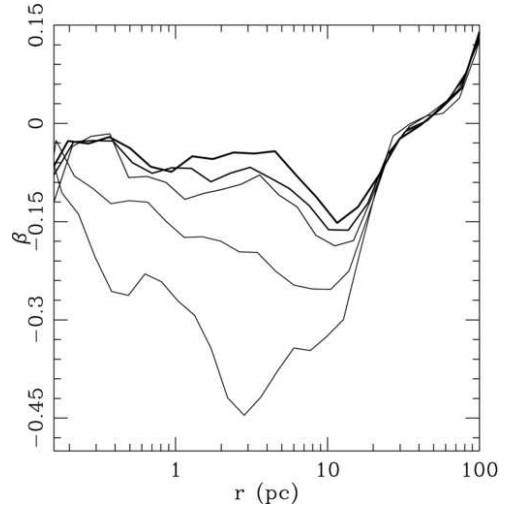


Figure 3: Evolution of the anisotropy parameter β during the post-merger phase. Times shown are the same as in Fig. 2; line thickness increases with time (from [1]).

are conveniently summarized as follows:

- The stellar system grows due to following merger events, and, after merging, shows a radial density profile that falls off as r^{-2} and a core of radius ~ 1 pc. These properties are similar to those observed in the MW NSC (see Fig. 1).
- The morphology of the NSC evolved during the infalls, from a strongly triaxial shape toward a more oblate/axisymmetric shape (see Fig. 2, where the triaxial parameter T is such that $T=0$ for oblate configurations, $T=1$ for prolate and $T=0.5$ for the ‘maximal’ triaxiality).

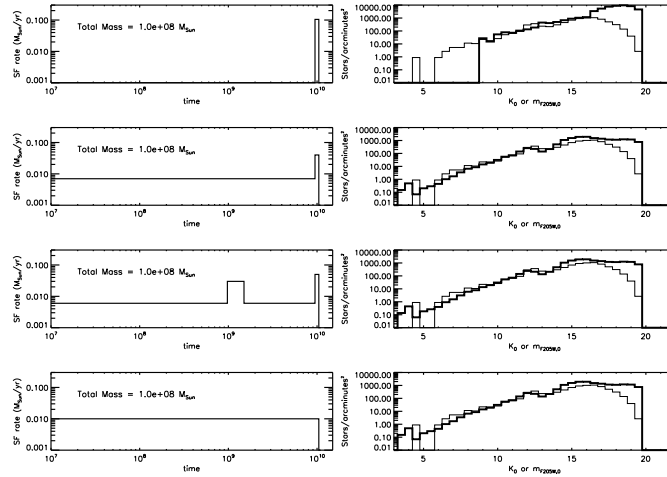


Figure 4: Left panels display star formation scenarios in which, from top to bottom: the entire mass is built up by ancient stars, brought into the Galactic center by infalling GCs; half of the mass is contributed by old stars and half due to continuous star formation; at the mass contributed by continuous star formation and GCs we add stars formed during a starburst episode occurred at ~ 1 Gyr; all the mass is contributed by continuous star formation. Right panels compare the observed LF of the Galactic center (light lines) with the LFs (heavy lines) resulting from the different star formation scenarios assuming solar metallicity and canonical mass-loss rates. Note that the data are much more than 50% incomplete for the faintest few bins. The models have not been scaled for mass, but rather have been rescaled along the vertical axis to match the number counts in the $K = 11.0$ bin (from [1]).

Kinematically, the final system is characterized by a mild tangential anisotropy within the inner 30 pc and a low degree of rotation (see Fig. 3).

- The effect of gravitational encounters on the evolution of the NSC was investigated continuing N -body integrations after the last inspiral event. The NSC core shrinks by roughly a factor of two in 10 Gyr as the stellar density starts evolving toward a Bahcall-Wolf cusp. This final core size is essentially identical to the size of the core observed at the center of the Milky Way. The density profile outside the core remained nearly unchanged during this evolution. Gravitational encounters also make the NSC to evolve toward spherical symmetry in configuration and velocity space.
- Since GCs are ancient objects with ages ~ 13 Gyr, in the merger model a large fraction of the NSC mass is expected to be in old stars. Using stellar population models, we showed that the observed LF of the MW NSC is consistent with a star formation history in which a large fraction (about 1/2) of the mass consists of old (~ 10 Gyr) stars and the remainder is from continuous star formation (see Fig. 4).

4. Conclusion and future developments

The migratory-merger model for the formation of dense stellar systems in the centers of elliptical galaxies and of superdense Nuclear Star Clusters in spirals has been applied to the case of the

Milky Way Nuclear Star Cluster, where a massive black hole dominates the innermost potential. Our N -body simulations show that a NSC with the same spatial and kinematical characteristics of the MW NSC is formed by inspiralling of a limited number ($\sim 10 - 15$) massive globular clusters in a time short respect to the age of the galaxy. Moreover, the observed luminosity function of the MW NSC is compatible with a younger population emebded in that coming from the merger object.

Relevant further steps forward in this topic are to be done. They include: i) more detailed merger simulation, using higher N sampling; ii) more extended and detailed *post* merger, secular, evolution; iii) a set of high precision simulations spanning a range of galactic central BH masses, to check whther more massive BHs may break the incoming GCs, halting the formation of a local dense stellar system, whose absence is noted in galaxies borrowing SMBHs at their center. Our group of computational astrophysics at the Dep. of Physics in Roma is presently dealing with these points, mainly by mean of our parallel, high precision and high performance N -body code (HiGPUs, see [5]) running on composite CPU+GPU platform.

References

- [1] Antonini, F., Capuzzo-Dolcetta, R., Mastrobuono-Battisti, A., Merritt, D. 2012, ApJ, 750, 111
- [2] Arca-Sedda, M., Capuzzo-Dolcetta, R. 2012, ASP Conf. Proc., vol. 453, p.337, ed. by R. Capuzzo-Dolcetta, M. Limongi, and A. Tornambè.
- [3] Böker, T. 2008, J. Phys. Conf. Ser., 131, 012043
- [4] Capuzzo-Dolcetta, R. 1993, ApJ, 415, 616
- [5] Capuzzo-Dolcetta, R., & Spera, M., & Punzo, D. 2013, JCP, 236, 580
- [6] Loose, H. H., Krügel, E., & Tutukov, A. 1982, A&A, 105, 342
- [7] Merritt, D. 2010, ApJ, 718, 739
- [8] Pesce, E., Capuzzo-Dolcetta, R., Vietri, M. 1992, MNRAS, 256, 368
- [9] Tremaine, S. D., Ostriker, J. P., & Spitzer, L., Jr. 1975, ApJ, 196, 407
- [10] Capuzzo-Dolcetta, R., & Vicari, A. 2005, MNRAS, 356, 899
- [11] Volonteri, M., Haardt, F., Gültekin, K. 2008, MNRAS, 384, 1387