

THEORY OF FORBUSH DECREASE IN A MAGNETIC CLOUD

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A model of Forbush-decrease in a magnetic cloud is presented. We calculate moments of the particle distribution function depending on time. The calculated results of cosmic ray intensity generally agree with the observed ones in events registered by ground based detectors. It is found that the magnetic flux rope is of great importance in dynamics of Forbush decrease.

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1. Introduction

Coronal mass ejections (CMEs) affect near-Earth space. They can include magnetic clouds (MCs), which have the strongest influence on Forbush decreases (FD). Theoretical studies of the spatial distribution of cosmic rays (CRs) are based on the solution of the diffusive transport equation. MC is presented as an infinitely extended cylinder with a constant radius moving in interplanetary space in [1], the expansion of the cylinder is taken into account in [2]. However, the cylindrical model has its shortcomings as spatial distribution of cosmic rays (CR) depends on magnetic flux rope (MFR) due to their high mobility. Thus, the toroidal model is more appropriate for these studies [3,4].

We present a MFR model in a MC, a method for calculating the CR distribution function in a MC, and the relations with its three moments. The calculation results and the measurements are compared.

2. MC model and calculation method

We suggest that a MC having the shape of a torus segment is located inside a CME with a fixed distribution of the radial velocity at the initial moment. The lateral boundaries of the segment will be referred to as ends. The initial MFR is described by the solution of Miller and Turner [5]. The subsequent MC motion is determined by the kinematic model, which describes the velocity conservation of the Lagrangian particles in a CME. The MFR during its motion is determined by the freezing-in condition represented as the conservation of the magnetic flux through the areas associated with the Lagrangian particle flow. The calculation of the particle distribution function is based on the Liouville theorem: the density of the particle distribution function retains along their trajectories, which are characteristics of the Boltzmann equation. We assume that: 1) Parker's solar wind is outside the MC; 2) the density of the distribution function is isotropic and homogeneous in the region ahead of MC.

The model MC is a segment of the magnetic cloud. Real MCs are connected to the Sun by means of magnetic field lines and the model MC describes their central parts. The MFR (the field direction changes from perpendicular on the surface to longitudinal at the center of the cross-section) determines the particle propagation. As a result, particles from the surrounding space fill only the near surface region of the cloud segment. A significant part of the particles comes to its ends. The particles can fill the near axis region when they return to the segment after redistribution along the field lines in the regions connecting the MC to the Sun. We introduce the loss factor of particles at the ends of $\alpha = (I_1 - I_2)/I_1$, which is a free parameter. Here I_1, I_2 are the particle fluxes coming to the ends and returning back to the segment. It is also assumed that the distribution function of the returning particles is isotropic and homogeneous over the cross-section of the end. The accepted assumptions are sufficient to calculate the particle distribution function.

In theory, the particle distribution function is calculated but in experimental studies the moments of the distribution function are determined. To obtain the relations between the distribution function and its moments, we write the relation

$$f(\vec{p}) = \sum_i n_i b_i + \sum_{i,j} n_i n_j f_{ij}, \quad (1)$$

where f is distribution function; $n_i = p_i/p$ is the unit vector of the particle motion direction; p_i, b_i are the momentum and vector components, respectively; f_{ij} is the symmetric 2nd-order

tensor. The components of the unit vector is defined by two angles. We determine the components of the vector and six components of the tensor by multiplying (1) by various trigonometric functions of these angles and integrating over the solid angle. For 2 moments we obtain

$$f_0 = \int f d\Omega / 4\pi = (f_{11} + f_{22} + f_{33})/3, \vec{A} = -(\vec{b}/f_0)100\%, \quad (2)$$

where f_0 is the distribution function density averaged over the solid angle (the zero moment), \vec{A} is anisotropy (the first moment), and $d\Omega$ is the differential of the solid angle.

Any symmetric 2nd-order tensor can be reduced to the diagonal form. For this purpose, we calculate the eigenvalues and eigenvectors of the tensor. The diagonal form of the tensor corresponds to the ellipsoidal form of the angular dependence of the distribution function referred to as the second moment. The ellipsoid shape depends on the ratio of the eigenvalues, with its orientation being determined by the eigenvectors.

3. Calculation results and discussion

We used the following parameters of the initial torus segment in the demonstration calculation: the torus cross-section radius is $0.043r_e$ (r_e is the astronomical unit), the torus axis is located in the solar equator plane at $0.457r_e$ from the Sun; angular width of the segment is $\pi/2$; the magnetic field intensity on the torus axis is 20 nT; MFR is SWN-type; the torus expansion during its radial motion is produced by a spherically symmetric CME, the velocities of the Lagrangian particles of which are distributed linearly along the radius from 500 to 400 km/s within the torus. MC is surrounded by the Parker's solar wind. The interaction between the MC and the solar wind is not taken into account. The parameters shown in the figures refer to the point with the coordinates of $x = 1r_e, y = z = 0$ in the heliocentric Cartesian coordinate system. The chosen point is located in the plane dividing the torus segment into equal parts. Time step is 1 hour. Kinetic energy of particles is $\varepsilon_k = 10GeV$. Figure 1 shows the magnetic field magnitude and its components in the MC in time. As can be seen from the figure, the chosen point is located inside the MC in the interval from 42 to 61 hours. The maximum value of the magnetic field is 20 nT, it did not change during the MC propagation.

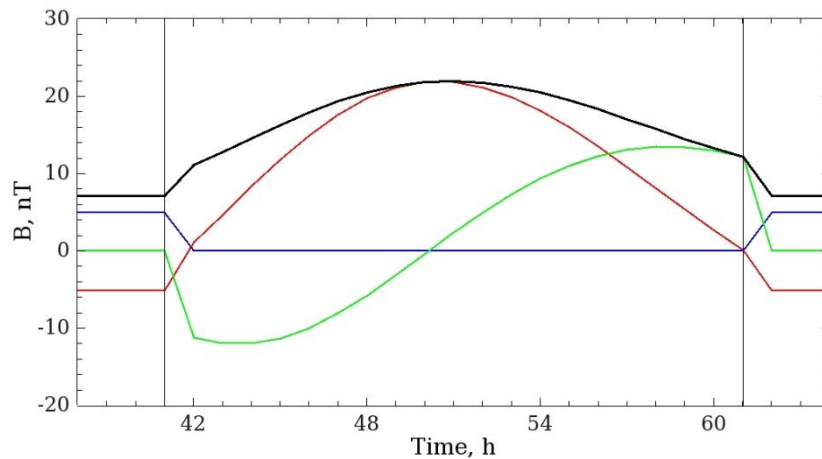


Figure 1: Magnetic field magnitude and its components: black line is vector magnitude, blue line is B_x , red line is B_y , green line is B_z . The vertical lines are MC boundaries.

Figure 2 shows the FD amplitude of $A_{FD} = ((f_0 - f_{00})/f_{00})100\%$, where f_{00} is the distribution function density outside the MC. The curves from top to bottom correspond to the values set of $\alpha = (0 - 5)\%$. The maximum value of A_{FD} changes from 0.2 to 2.5%. Thus, the greater the loss factor at the segment ends, the greater A_{FD} .

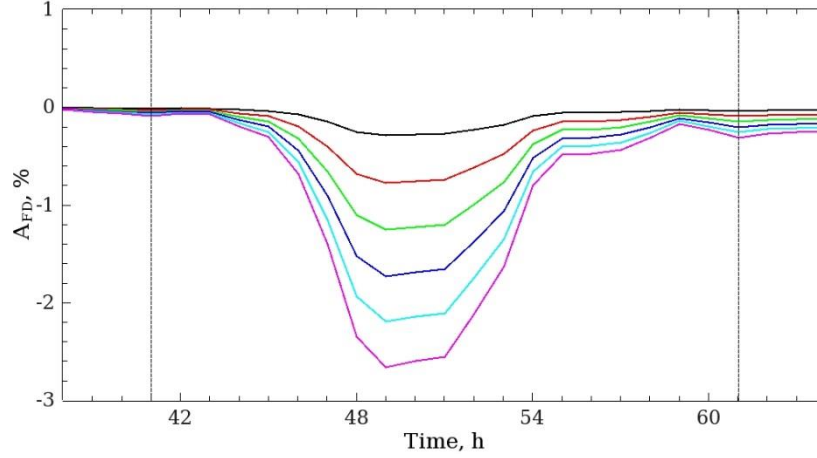


Figure 2: Amplitude of FD in time. The curves from top to bottom correspond to $\alpha = (0 - 5)\%$. The vertical lines are MC boundaries.

Figures 3, 4 show the anisotropy components with $\alpha = 5\%$ as a time function for two MFR types: SWN type in Figure 3 and SEN type in Figure 4. Each component has a feature in temporal behavior. Component A_x is the smallest, components A_y, A_z change their sign two and one time crossing the MC, respectively. One can see stable dependences in the behavior of the anisotropy components. In the region with the maximum values of component B_y (interval (48-55) hours according to Figure 1), A_y has a minus sign. For component A_z , the MC crossing time can be divided into two intervals: the first interval starts from the entrance into the MC and ends at its center; the second one starts from the center and ends at the exit from the MC. The component sign in the first interval depends on the sign of component B_y : $A_z > 0$ when $B_y > 0$ and vice versa. The boundary between the intervals coincides with the sign change of component B_z , i.e. in the vicinity of the magnetic cloud center. For NWS and NES types A_y has a plus sign. Sign of component A_z in the first interval depends on the sign of component B_y : $A_z > 0$ when $B_y > 0$ and vice versa, as well as for SWN and SEN types. These relations confirm the possibility to determine the magnetic field properties based on measuring the CR intensity by ground-based detectors [6].

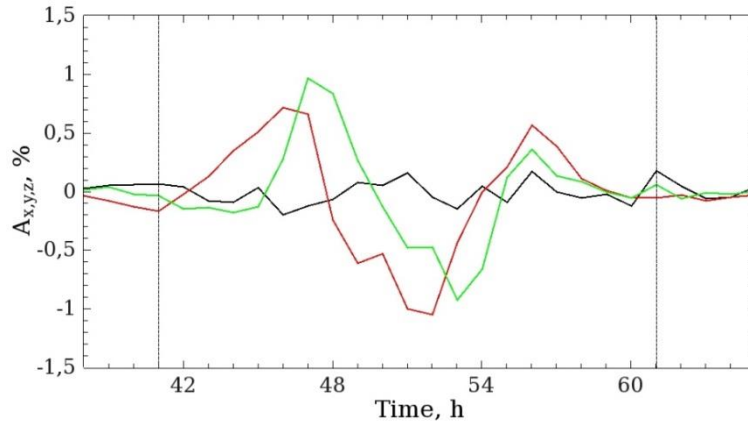


Figure 3: Anisotropy components with $\alpha = 5\%$ in time: black line is component A_x , red line is A_y , green line is A_z . SWN type of the MC.

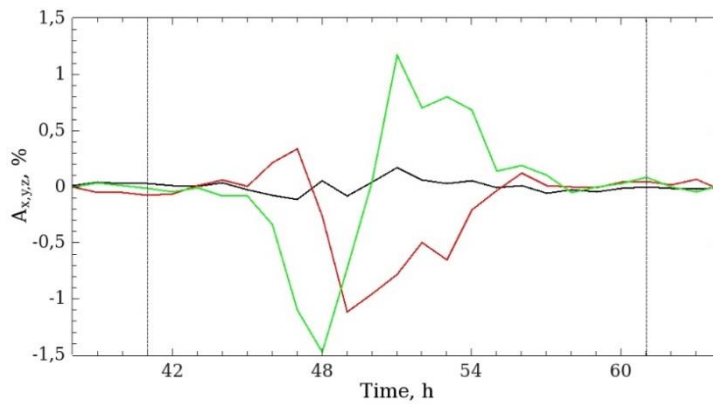


Figure 4: Anisotropy components with $\alpha = 5\%$ in time: black line is component A_x , red line is A_y , green line is A_z . SEN type of the MC.

The anisotropy magnitude depends on the loss factor α at the ends of the magnetic cloud segment. Figure 5 shows the anisotropy components depending on time for different α . The components with $\alpha=5$ coincide with the model ones shown in Figure 3. As can be seen, the behavior of the anisotropy components does not depend on α in time, however, the anisotropy amplitude shows a dependence: the smaller α , the less the anisotropy, as well as the FD amplitude.

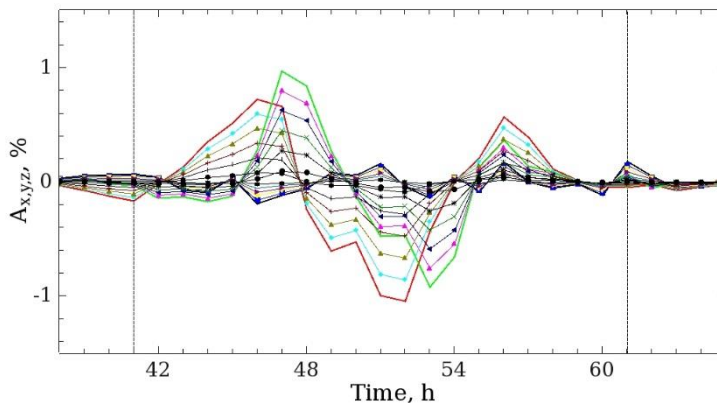


Figure 5: Anisotropy components with $\alpha = (0 - 5)\%$ in time. SWN type of the MC.

Figure 6 illustrates the equatorial component of anisotropy in the form of vector sum of hourly values of A_x , A_y , which are the anisotropy components. As can be seen from the Figure 6: 1) the anisotropy significantly increases in the MC (43-57) hours; 2) component A_y is significantly higher than component A_x ; 3) component A_y increases within (42-47) hours and decreases within (47-54) hours; 4) the anisotropy vector rotates: the rotation is anticlockwise in the interval during (45-51) hours and clockwise during (51-56) hours.

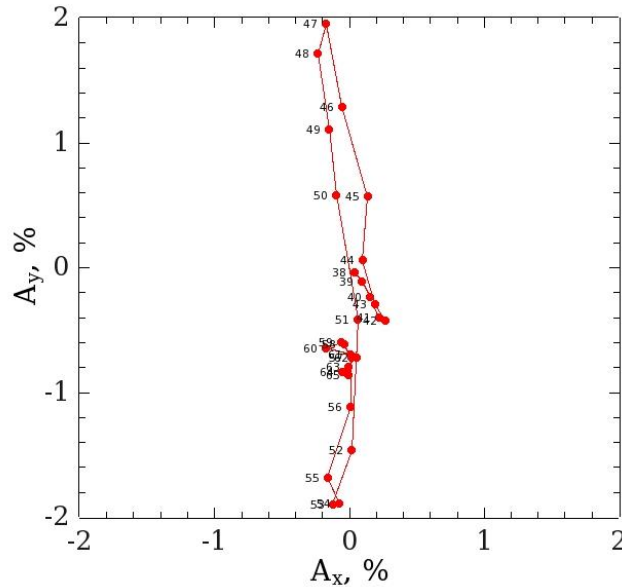


Figure 6: Equatorial component $A_{x,y}$ of the anisotropy. SWN type of the MC.

Figures 7,8 show the eigenvalues and the orientation angles of the second moment of the CR distribution function. As can be seen from Figure 7, the difference between two eigenvalues is $\sim 0.3\%$, and their difference from the third eigenvalue is $\sim 3\%$ in the interval (46-54) hours. Outside the interval, the eigenvalues coincide with the accuracy of $\sim 0.3\%$. Thus, the second moment of the CR distribution function in the interval (46-54) hours represents a biaxial ellipsoid, in which the smallest of the axes differs from the other two by $\sim 3\%$. Outside this time interval, the second moment represents the sphere with the accuracy of $\sim 0.3\%$. Comparing Figures 2 and 7, it is clear that the ellipsoid of the second moment coincides with the FD in time.

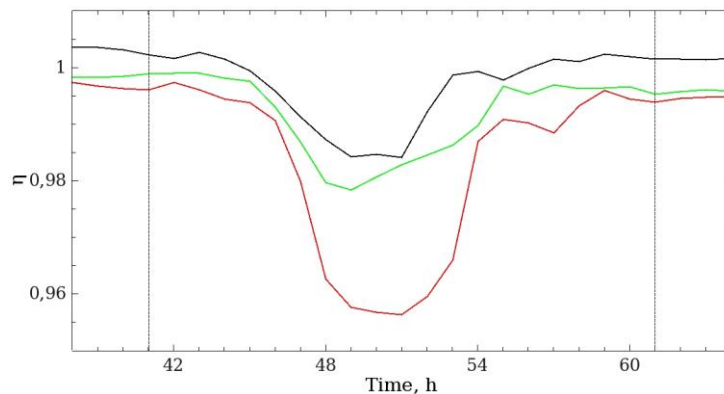


Figure 7: Eigenvalues of the second moment of the CR distribution function in time.

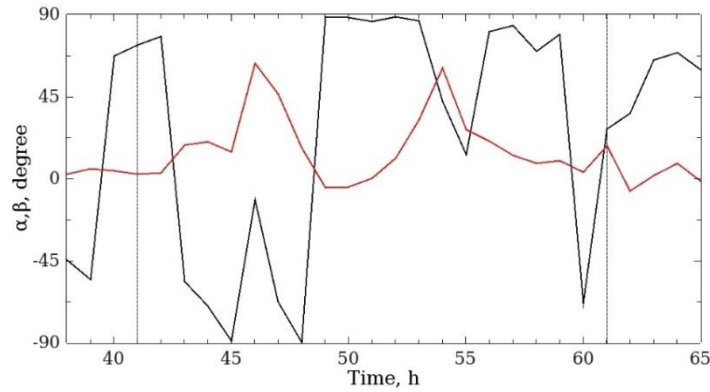


Figure 8: Orientation angles of the second moment of the CR distribution function in time. Black curve is course angle and red one is tangage angle.

The orientation of the ellipsoid in space is determined by 3 Euler angles: course, tangage and bank. The orientation of a biaxial ellipsoid is determined by 2 angles: α is course angle - the angle between the projection of the smallest axis on the plane XOY and the X axis; β is tangage angle - the angle between this projection and the Z axis.

As can be seen from Figure 8, in the interval (48-54) hour, $\alpha \approx \pi/2$ (as for the ellipsoid axis $\alpha = -\pi/2 = \pi/2$) and β monotonically changes from $\pi/4$ to 0 and back to $\pi/4$.

The calculation of the CR distribution function with 60 GeV energy in the MC with similar parameters showed the absence of FD. It reveals that the FD amplitude depends on the magnetic field strength and the MC cross-section radius. The presented results are obtained for a MC, whose axis is located in the solar equator plane. Comparing the calculation results with the measurements, it is necessary to take into account the trajectory of the MC crossing, on which the time dynamics of the moments of the CR distribution function significantly depends.

The work [7] presents the results of the analysis of the CR density and anisotropy with the rigidity of 10 GV (9 GeV) in 99 events of the 23rd and 24th solar cycles, in which the presence of MCs is confirmed.

In the studied sample, there is a great diversity in: motion velocity, the gradient of the flow velocity within the MC, the distribution and magnetic field magnitude in the MC. The time dynamics of the CR density and anisotropy of is also diverse. In this sample, the characteristic behavior of density and anisotropy in MCs is presented: 1. the maximum value of A_{FD} is within (1-12)%; 2. the density drop is sharp, and the recovery is more gradual; 3. the anisotropy changes significantly at the entrance and/or exit from the MC; 4. monotonic change of the anisotropy is often observed inside the MC; 5. a rotation of the equatorial component $A_{x,y}$ is possible; 6. the north-south component (A_z) often changes its sign in the vicinity of the MC center. As can be seen from the results presented in Figures 2-6, the model calculations generally coincide with the observed behavior of the density and anisotropy in the MC.

4. Conclusions

1. We show a model of a dynamic MC located in an inhomogeneous matter flow. A method for calculating the MFR in the MC initially being a torus segment is presented and applied.

2. A method for calculating the particle distribution function in the MC using particle trajectories is presented and applied.

3. The relations between the particle distribution function and its 3 moments are obtained.

4. The influence of the regions connecting MC with the Sun on the CR distribution in the MC is determined. This effect is due to the MFR.

5. Time dynamics of 3 moments of the particle distribution function in the MC is calculated. The dynamics of the CR density and anisotropy generally agree with the measurements.

6. The found relations between the dynamics of the distribution function moments and the MFR confirm the possibility to determine the magnetic field properties based on measuring the CR intensity by ground-based detectors.

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