

## Estimating the Depth of Shower Maximum using the Surface Detectors of the Pierre Auger Observatory

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**C. J. Todero Peixoto\*** for the Pierre Auger Collaboration<sup>†</sup>

*Escola de Engenharia de Lorena, Universidade de São Paulo*

*E-mail: [toderocj@usp.br](mailto:toderocj@usp.br)*

The Pierre Auger Collaboration has developed a method for estimating the depth of shower maximum ( $X_{\max}$ ) using the risetime of the signals recorded by the array of water-Cherenkov detectors. The average value of  $X_{\max}$  is determined for energies above 3 EeV, using air showers with zenith angles up to  $60^\circ$ . The work described here considerably extends that reported previously with an increase in the total number of events of a factor 2.1. In total 125,005 events are used, of which 237 are above 50 EeV. The data set is approximately a factor 30 larger than what is available from fluorescence measurements for energies above 3 EeV. The results reinforce previous conclusions: the mean mass of high-energy cosmic rays begins to rise above 3 EeV and extends that finding to higher energies with greater precision. Additionally, we find strong discrepancies when we make a comparison between observations and predictions from simulations with values which are obtained from measurements of the risetime of the recorded signals.

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\*Speaker.

<sup>†</sup>for collaboration list see PoS(ICRC2019)1177

## 1. Introduction

Knowledge of the mass composition of cosmic rays would provide us with crucial information to discriminate between different classes of production, acceleration and propagation models, and to understand the origin of the observed flux suppression; it would allow us to search for a stream of protons at higher energies, thus enhancing the study of small-scale anisotropy. It could further improve previous studies of particle physics at 10 EeV and extend them to energies as high as  $\sqrt{s} = 200$  TeV.

Using data from the Auger Observatory several methods that enable mass discrimination with ground based detectors have been developed. Of the several parameters studied, those that presented some relevance were  $X_{\max}$  [1] (from measurements of the fluorescence detector, FD),  $X_{\max}^{\mu}$  [2], Risetime Asymmetries [3] and the Delta Method [4,5], the latter three using the signals from the surface detector, SD. The  $X_{\max}$  parameter has the best resolution for mass composition studies, but the FD only operates during nights without moonlight, only 13% of the operating time of the SD. With the Delta method, we can use the signals from the SD, which has a nearly 100% duty cycle, to calculate a parameter,  $\Delta_S$ , and to convert this to an estimation of the depth of the shower maximum,  $X_{\max}^{\text{Delta}}$ .

This work is based on the Delta Method [4], which was originally used with SD data up to zenith angles  $\sim 45^\circ$ . We now extend the method to events up to  $60^\circ$ . This yields an increase of about 50% in the number of events, thus reducing the statistical uncertainty associated with mass composition inferences. We have also extended the sample by including three additional years of data. The large data set from the 750 m array has not been updated.

## 2. Data Selection

This work uses a data sample from the 1500m array collected from January, 2004 to August, 2018. We consider events with zenith angles below  $60^\circ$ , and with energies above 3 EeV to guarantee full trigger efficiency.

We also impose conditions at the level of single WCDs (Water-Cherenkov Detectors): the total signal must be above 5 VEM and the distance to the shower core must be in the interval [300, 1400] m for the 1500m array. For events with energies  $\log_{10}(E/\text{eV}) > 19.8$ , we extend the distance range up to 2000m.

After applying these selection criteria, the number of events used in what follows is 125,005.

## 3. The Delta Method

The Delta Method is explained in Figure 1(Left). We characterize every event with a single parameter  $\Delta_S$ , which is a function of the risetimes<sup>1</sup> recorded for this event. We

<sup>1</sup>This parameter is the time for the integrated signal to increase from 10% to 50% of the final magnitude of the integrated signal.

have determined relationships (dubbed as *benchmarks*) that describe the risetimes as a function of the distance and zenith angle in a narrow energy range. Risetimes for each WCD are compared to the predicted *benchmark* value, in units of the accuracy with which risetimes are determined, giving a parameter  $\Delta_i$ .  $\Delta_s$  is the average of the  $\Delta_i$  values over the  $N$  selected WCDs. An example of the dependence of the risetime with the distance to the core is shown in Figure 1(Right).

For each PMT in a detector there are two FADC (Flash Analog-to-Digital Converter) channels: a high-gain and a low-gain channel (used when the high-gain (HG) channel has saturated). Fits to the data are made using the function of equation 3.1 (with  $A$  and  $B$  free parameters). With  $A$  and  $B$  known, a fit to the points from the HG channel is made (shown by grey points in Figure 1(Right)) with  $N$  as a free parameter (equation 3.2). In this way the small differences between the measurements with the two channels are dealt with.

These three parameters are parameterized as a function of  $\sec \theta$  in the narrow *benchmark* energy bin. The energy bin used is:  $\log_{10}(E/eV) = 19.1-19.2$ .

$$t_{1/2}^{low-gain} = 40 \text{ ns} + \sqrt{A(\theta)^2 + B(\theta) \cdot r^2} - A(\theta) \quad (3.1)$$

$$t_{1/2}^{high-gain} = 40 \text{ ns} + N(\theta) \cdot \left( \sqrt{A(\theta)^2 + B(\theta) \cdot r^2} - A(\theta) \right) \quad (3.2)$$

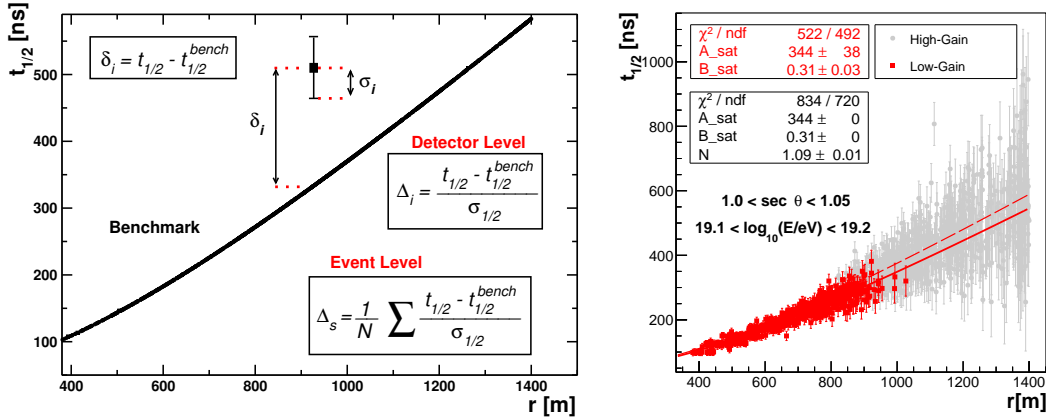


Figure 1: (Left) Description of the Delta Method. (Right) Risetime measurements as a function of the distance to the core for data from the 1500m array in the energy interval  $19.1 < \log_{10}(E/eV) < 19.2$ .

The Delta Method procedures are fully described in [4,5].

#### 4. Extending to angles up to 60 degrees

Once the functional form of  $A(\theta)$ ,  $B(\theta)$  and  $N(\theta)$  are known, the mean value of the  $\Delta_s$  distribution must be, by definition, compatible with zero for the *benchmark* energy bin, as shown in the Figure 2(a) for events with  $\sec \theta < 1.45$ . For more inclined events, the parameter  $A$  falls rapidly to zero, so that a specific parameterization is needed for this

set of events. The result of this new parameterization is shown in the Figure 2(a). As expected, the mean value of the distribution for more inclined events is compatible with zero.

Merging these datasets ( $\sec \theta < 1.45$  and  $\sec \theta > 1.45$  in Figure 2) we see that the new  $\langle \Delta_S \rangle$ , Figure 3(a), is consistent with zero. The  $\langle \Delta_S \rangle$  as a function of  $\sec \theta$  is shown in the Figure 3(b). In the *benchmark* region,  $\Delta_S$  shows no dependence on the angle.

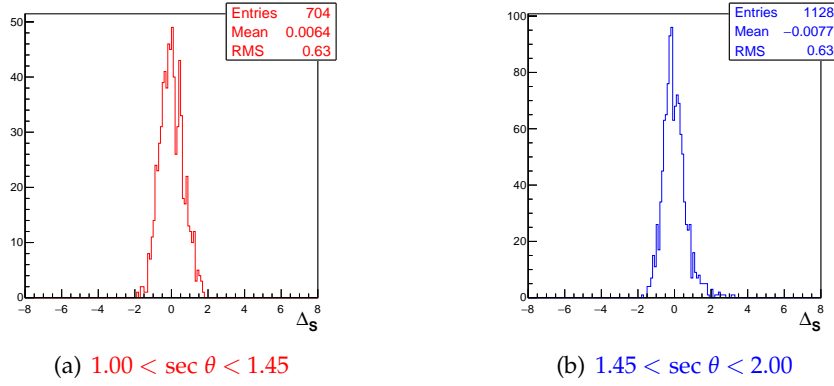


Figure 2:  $\Delta_S$  values for two zenith angle ranges in the energy bin:  $19.1 < \log_{10}(E/\text{eV}) < 19.2$ .

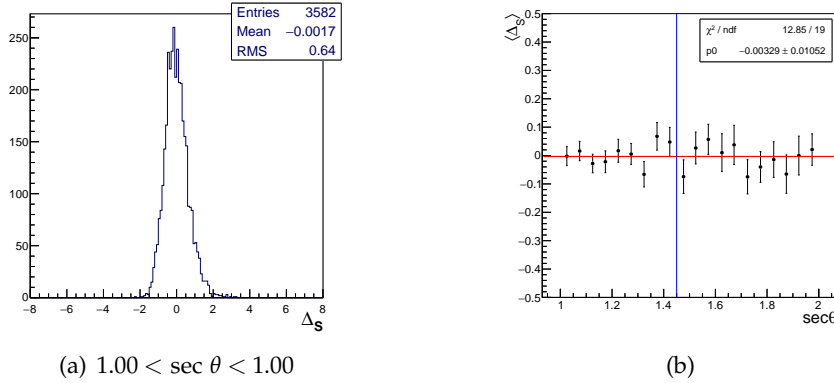


Figure 3: (a) Distribution of  $\Delta_S$  values for the whole angular range considered in this work. (b) Dependence of  $\langle \Delta_S \rangle$  with the zenith angle.

## 5. $\langle \Delta_S \rangle$ as function of the energy and the evolution for $\langle \ln A \rangle$ with energy

In Figure 4(Left) the evolution of  $\langle \Delta_S \rangle$  as a function of energy is shown. In Figure 4(Right) a comparison with hadronic models is made. The measurements are well within the predictions from simulations and they favour a composition dominated by heavier elements as the energy increases. In Figure 5 good agreement between the new and the previously published data is shown. In total we have more than doubled the number of events. For  $\log_{10}(E/\text{eV}) > 19.8$ , the data set has been increased by a factor 2.1 because of the larger distance range adopted in the signal selection. Our results confirm, with a bigger statistical sample, that as the energy increases, the composition of the flux is dominated by elements heavier than protons.

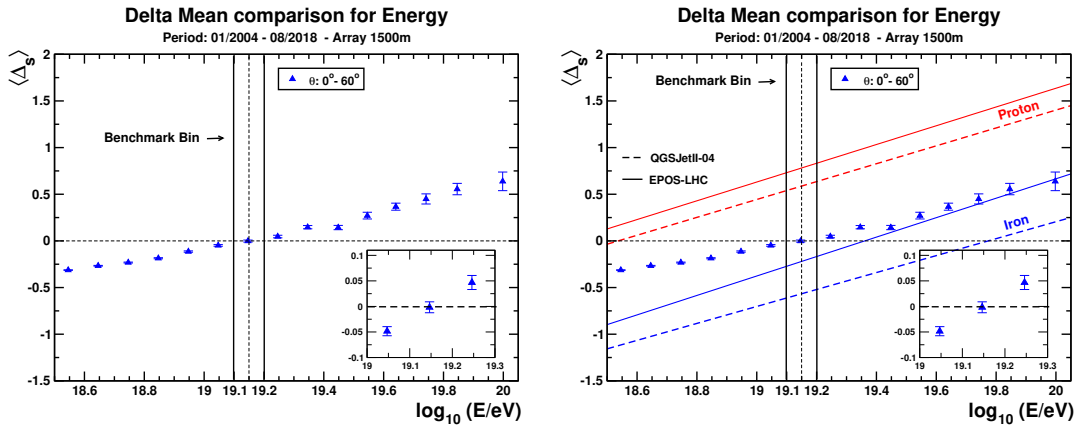


Figure 4: (Left)  $\langle \Delta_S \rangle$  as function of the energy for the new dataset. (Right) Comparison of the evolution of  $\langle \Delta_S \rangle$  as a function of the energy for data and simulations.

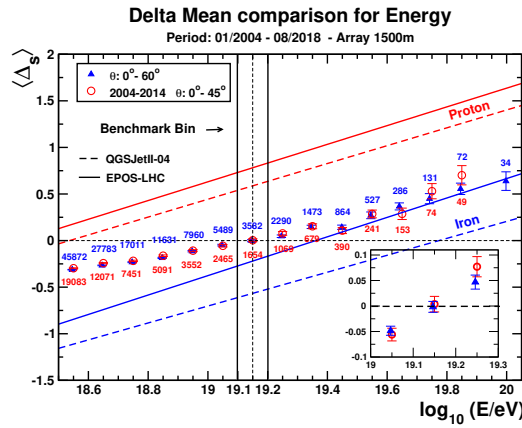
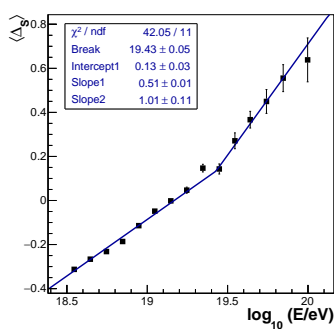
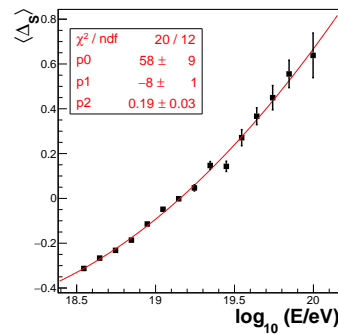


Figure 5: Comparison with previously published results. The numbers indicate the amount of events in each energy bin.



(a) Broken Line.



(b) Second degree polynomial.

Figure 6: Fits for  $\langle \Delta_S \rangle$  as function of the energy.

In Figure 6, we fit our  $\langle \Delta_S \rangle$  measurements to different functional forms. The fit to a straight line is poor ( $\chi^2 = 75$  for 13 degrees of freedom). When a broken line is considered the fit is also poor ( $\chi^2/\text{dof} = 42/11$ ). A better fit is obtained ( $\chi^2/\text{dof} = 20/12$ ) when our measurements are fit to a second degree polynomial. When compared to model predictions, this result clearly indicates that the rate of change of mass composition varies with energy.

Although it is not possible to compare directly the fluorescence and surface detector data, it is possible to transform the  $X_{\text{max}}$  and the  $\Delta_S$  values to values of  $\langle \ln A \rangle$ . In this work we used the proton and iron primaries for the hadronic models QGSJetII-04 [6] and EPOS-LHC [7].

Figure 7 shows that the QGSJetII-04 and EPOS-LHC models do not describe data adequately. One of the most prominent reasons for this disagreement comes from the difference observed when comparing the number of predicted muons in models with the measurements made with the Pierre Auger data [8].

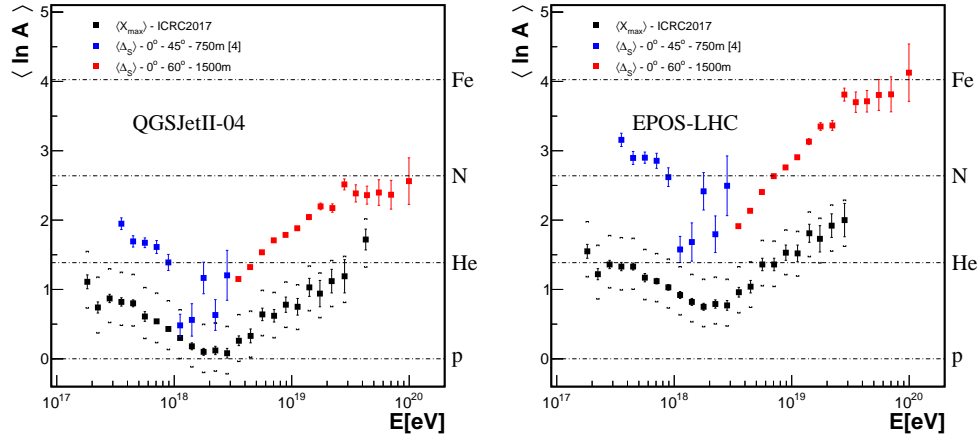


Figure 7: Evolution of the average mass with energy for measurements done with the FD and the SD.

## 6. Cross Calibration between $\Delta_S$ , $X_{\text{max}}^{\text{FD}}$ and Energy

We now address the correlation of  $\Delta_S$  with  $X_{\text{max}}$  using hybrid events. We would not expect a 1:1 correlation between these parameters because the position of  $X_{\text{max}}$  is dominated by the electromagnetic component, whereas  $\Delta_S$  is dependent on a muon/electromagnetic mix. The values of  $\Delta_S$  can be calibrated with  $\langle X_{\text{max}}^{\text{FD}} \rangle$  providing a correlation from which a new parameter  $\langle X_{\text{max}}^{\text{Delta}} \rangle$  can be calculated. These events have been taken from the FD data set discussed in [9] and are show in Figure 8. There are 2126 events available for calibration. The selected samples of events are unbiased.

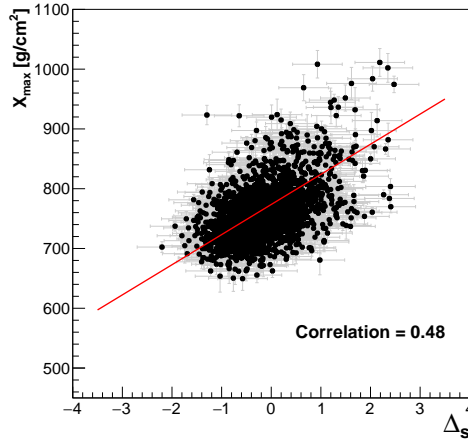
A linear surface (equation 6.1) was adjusted and the parameters obtained were used to calculate  $\langle X_{\text{max}}^{\text{Delta}} \rangle$ . The values of the parameters  $a$ ,  $b$  and  $c$  are listed in Table 1.

$$\langle X_{\text{max}}^{\text{Delta}} \rangle = a + b \cdot \langle \Delta_S \rangle + c \cdot \log_{10}(E_{\text{SD}}/\text{eV}) \quad (6.1)$$

Array	$a(\text{g}/\text{cm}^2)$	$b(\text{g}/\text{cm}^2)$	$c(\text{g}/\text{cm}^2)$
1500m	$733 \pm 38$	$34 \pm 12$	$2 \pm 1$

Table 1: Values for  $a$ ,  $b$  and  $c$  for equation 6.1.

The  $\langle X_{\text{max}}^{\text{Delta}} \rangle$  and  $\langle X_{\text{max}}^{\text{FD}} \rangle$  vs energy are shown in Figure 9 together with the number of events (FD and SD) for each energy bin. This analysis confirms with greater precision that, as the energy increases, intermediate/heavy nuclei are preponderant in the flux composition of ultra-high energy cosmic rays.

Figure 8: Correlation between  $X_{\text{max}}$  and  $\Delta_s$  for hybrid events.

## 7. Conclusions

The Delta Method uses the information provided by the risetimes of the signals recorded by the WCD of the Surface Detector Array. With it, we can probe the validity of the assumptions contained in state-of-the-art packages that model hadronic interactions at ultra-high energies. Likewise, we can perform mass composition studies with the largest data sample available. This work extends the applicability of the Delta Method up to zenith angles of  $60^\circ$  and adds three more years of data to previously published results. We confirm, with an improved precision, that: i) the rate of change of mass composition varies in the energy range covered by our analysis and, that intermediate/heavy nuclei are predominant as the energy increases; ii) simulations do not reproduce correctly the physics of extensive air showers.

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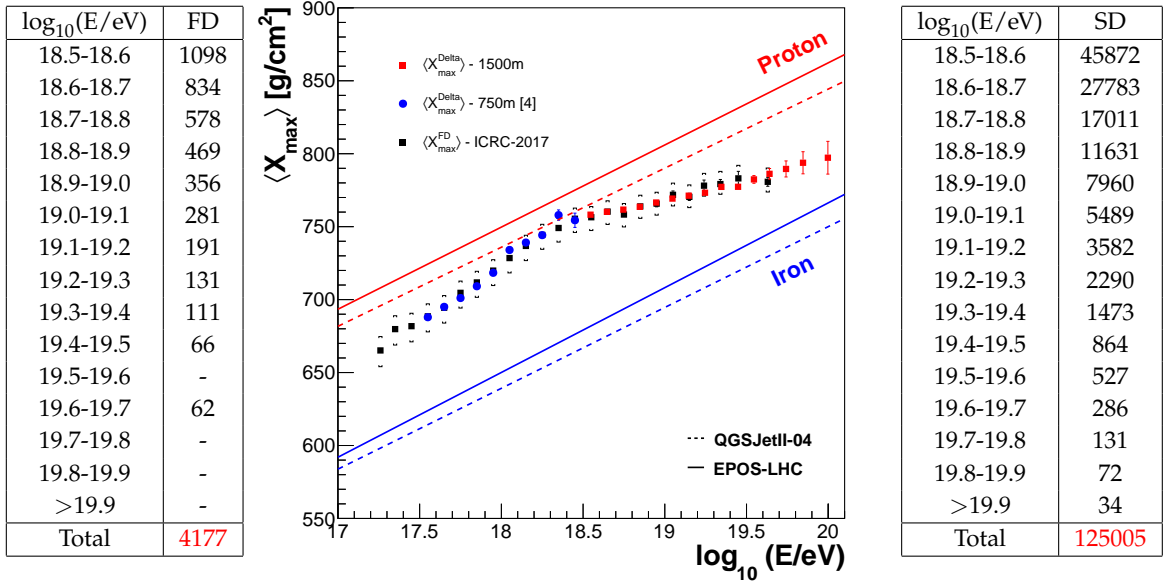


Figure 9: Evolution of  $\langle X_{\max}^{\text{Delta}} \rangle$  with energy. The SD measurements are compared with the  $\langle X_{\max} \rangle$  measurements done with the FD. The figures in the table correspond to the number of events per energy bin. They are restricted to the energy interval where the Delta method has been updated.

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