



The Application of 20 inch MCP-PMT In LHAASO-WCDA

Xiaohao You,^{a,*} Bo Gao,^a MingJun Chen,^a HuiCai Li,^a Cheng Liu^a and Kai Li^a on behalf of the LHAASO Collaboration

(a complete list of authors can be found at the end of the proceedings)

^aInstitute of High Energy Physics, 19B Yuquan Road, Shijingshan District, Beijing, China E-mail: youxiaohao@ihep.ac.cn

In the Large High Altitude Air Shower Observatory (LHAASO), the main physics objective of the Water Cherenkov Detector Array (WCDA) is to survey the gamma-ray sky continuously in the energy range from 100 GeV to PeV. The Water Cherenkov detector array, covering an area of about 78,000 m^2 area, is constituted by 3120 detector units divided into 3 separate arrays. In the second and third array are installed 2220 20" PMTs instead of the 8 PMT used in the first $150 \times 150m^2$ array. This type of PMT has large sensitive area, high quantum efficiency (QE), and large peak-to-valley (P/V) ratio for single photoelectron detection. In this work, we will report on the application of 20 inch MCP-PMT at LHAASO-WCDA.

37th International Cosmic Ray Conference (ICRC 2021) July 12th – 23rd, 2021 Online – Berlin, Germany

*Presenter

[©] Copyright owned by the author(s) under the terms of the Creative Commons Attribution-NonCommercial-NoDerivatives 4.0 International License (CC BY-NC-ND 4.0).

Xiaohao You

1. Introduction

The Water Cherenkov Detector Array, covering an area of about 78, $000m^2$ area, is constituted by 3,120 detector units divided into 3 separate arrays. Every array is a single water pond with 4.5m depth. Two of them with an effective area of $150 \times 150m^2$ contain 900 detector units each. The third array(WCDA-3) with an area of $300 \times 110 \text{ m}^2$ contains 1,320 detector units. Each detector unit is divided in $5 \times 5m^2$, separated by black plastic curtains vertically hung in the water to isolate the scattered light. A pair of 8" and 1.5" PMTs in each unit of WCDA-1, while a pair of 20" and 3" PMTs in WCDA-2 and WCDA-3.[1].

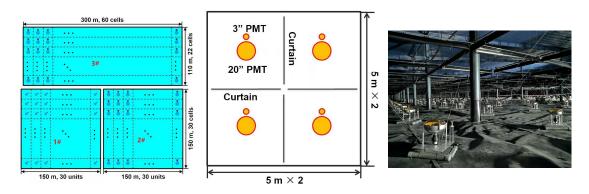


Figure 1: Schematic of LHAASO-WCDA. Left: full picture of the LHAASO WCDA, where are shown the three pools. Middle: Scheme of the arrangement of the PMTs in WCDA-3. Right: a picture of the inside of WCDA-3.

Table 1: Re	quirements of	of 20-inch PMTs
-------------	---------------	-----------------

No.	Item	PMT Specifications
1	Diameter of cathode	20 in, ellipsoid
2	Withstand Pressure	>3 atm
3	Working Voltage	Gain@ 5×10^{6} , <2000V
4	Working current	<300µA
5	Time resolution	<2ns
6	Peak-to-valley ratio	>2
7	Response curve	300nm-650nm
8	TTS	<7ns@5×10 ⁶
9	CTTD	<4ns@5×10 ⁶
10	Dark noise	<25 kHz@5 $\times10^{6}$, threshold = 1/3PE, @T = 25°C
11	Nonlinearity	$\geq 1800PE(\pm 10\%)$

2. The 20 inch MCP-PMT

The 20 inch MCP-PMT is manufactured by North Night Vision Technology Co., Ltd(NNVT) at Nanjing, China. The newly developed 20 inch PMT shown in Fig2, uses micro-channel-plate

(MCP) instead of the traditional dynodes enabling a better energy resolution and good detector response. It consists of bialkali photocathode, a focusing electrode, a MCP, and an anode. The distance between the photocathode and MCP is nearly 300 mm. Time resolution of the 20-inch MCP-PMT predominantly depends on the electrical field distribution between the photocathode and the MCP, therefore, a lotus-like focusing electrode was designed to reduse transit time spread(TTS) to 5.8 ns(FWHM)[2].

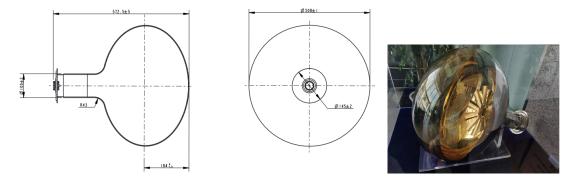


Figure 2: Dimension of 20 inch MCP-PMT.

The average dark noise of the 20-inch MCP-PMT was ~ 16 kHz at $\sim 5^{\circ}$ C in WCDA-3 during dry run mode, and it increased to ~ 51 kHz once filled with 4.5 meters of water, as shown in Fig. 3.

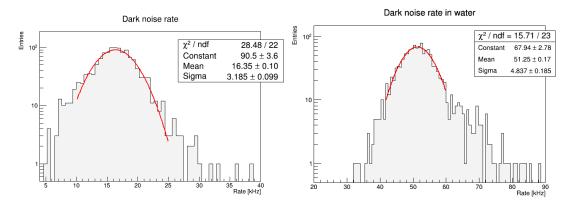
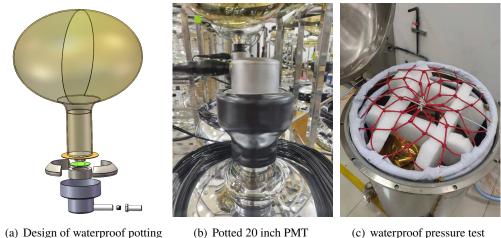


Figure 3: Left: dark noise rate without water, Right: dark noise rate in 4.5 meters of water.

3. PMT waterproof potting

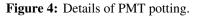
The 20 inch MCP-PMTs in LHAASO-WCDA had been anchored at 4.5m water depth and will operate at least for 10 years. Waterproof potting of PMT is one of the keys for system reliability. Based from JUNO's¹design, waterproof potting was designed to optimize signal quality and stability under water pressure. The potted detector is composed of PMT, HV divider and 30 meters cable (Fig.4(b)). Fig.4(a) show the structure details of waterproof potting.

¹The Jiangmen Underground Neutrino Observatory



(a) Design of waterproof potting

(b) Potted 20 inch PMT



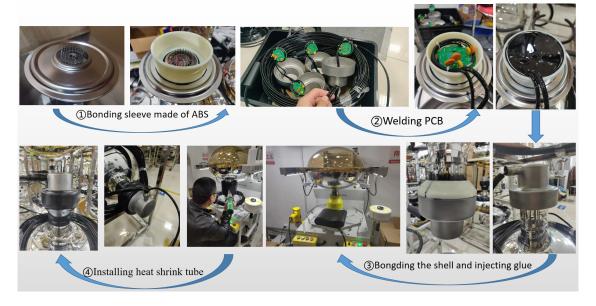


Figure 5: PMT potting procedure

Fig.5 shows the complete PMT waterproof potting procedure. To avoid damaging the PMT, the heat shrinkable tube is installed within 2 minutes in step 4.

Test Result of 20 inch MCP-PMT 4.

Each PMT had been tested at least three times before potting. In order to ensure the potting process didn't damage the PMT, all the 20-inch PMTs will be tested individually again before delivery to LHAASO site. The PMT test system includes two parts: cathode test system and anode test system.

- The cathode test system is mainly used to test the cathode sensitivity, quantum efficiency, and cathode non-uniformity of the PMT;
- The anode test system is mainly used to test the PMT single photoelectron spectrum, gain, working high voltage, single photoelectron spectrum, peak-to-valley ratio, and energy resolution, such as dark noise rate, transit time dispersion, rise time, fall time, response time, pre-pulse ratio, post-pulse ratio, dynamic range, etc.

All parameters meet LHAASO-WCDA's requirements (Tab 1). Part of test result are shown at Fig 6, where the mean working voltage $1777V(Gain@5 \times 10^6)$. The transit time Spread below 7 ns and the peak-to-Valley ratio greater than 2.

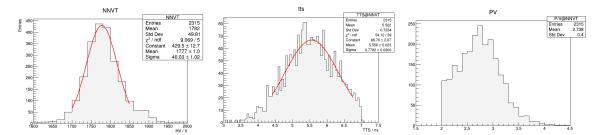


Figure 6: Part of test result at NNVT. Left: the distribution of working voltage, Middle: transit time spread, Right: peak-to-Valley ratio.

5. The Permalloy Geomagnetic shield

Due to the large size of 20 inch MCP PMT, the geomagnetic field have big influence on the performance of PMT, including time response, charge resolution and detect efficiency. A Permalloy shielding had been installed on all 20 inch MCP-PMT, as shown at Fig.7. It is a nickel-iron magnetic alloy, with about 80% nickel and 20% iron content. Permalloy can change the direction of the magnetic field because of its higher permeability compared to ordinary steel. Since the magnetic field is guided through a low magnetic resistance path, it ensures that the PMT is not affected by the geomagnetic field.

As shown in Fig.8, the time resolution performance of the PMT with magnetic shield is significantly better than that of the PMT without shield.

6. Summary

The 20-inch MCP-PMTs are working well at LHAASO-WCDA and the waterproof potting failure rate is less than 1% as of June 31,2021. LHASSO-WCDA is taking data and the total array results will be published in October 2021 on current schedule.

7. Acknowledgements

The authors would like to thank all staff members who work at the LHAASO site above 4400 meters above sea level year-round to maintain the detector and keep the electrical power supply



Figure 7: Left: full view of magnetic shield-Permalloy installed on 20 inch MCP-PMT. Right: top view of PMT.

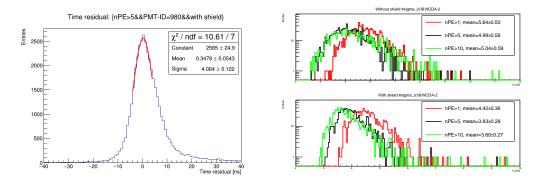


Figure 8: The difference of time residual between PMT with magnetic shield and not. Left: time residual distribution of a PMT with shield when number of photoelectrons(nPE) is great than $5(\sigma_r=4.08\text{ns})$. Right: Comparison of σ_r of time residual in 900 MCP-PMTs with/without magnetic shield. The red line represents the distribution of σ_r at nPE>1 ($E(\sigma_{without-shield})=5.64\text{ns}$, $E(\sigma_{with-shield})=4.42\text{ns}$), the black line represents the distribution of σ_t at nPE>5 ($E(\sigma_{without-shield})=5.00$ ns, $E(\sigma_{with-shield})=3.82\text{ns}$), the green line represents the distribution of σ_r at nPE>10 ($E(\sigma_{without-shield})=5.02\text{ns}$, $E(\sigma_{with-shield})=3.60\text{ns}$).

and other components of the experiment operating smoothly. We are grateful to the Chengdu Management Committee of Tianfu New Area for their constant financial support of research with LHAASO data.

References

[1] Cao Zhen, Chen Ming-jun, Chen Song-zhan, Hu Hong-bo, Liu Cheng, Liu Ye, Ma Ling-ling, Ma Xin-hua, Sheng Xiang-dong, Wu Han-rong, Xiao Gang, Yao Zhi-guo, Yin Li-qiao, Zha Min, and Zhang Shou-shan. Introduction to large high altitude air shower observatory. *Chinese Astronomy and Astrophysics*, 43(4):457–478, 2019. [2] Ling Ren, Jianning Sun, Shuguang Si, Xingchao Wang, Muchun Jin, Guorui Huang, Haiyang Xu, Congjie Wang, Wei Hou, Liang Wang, Zhen Jin, Fangjian Qiao, Xiangbiao Qiu, Mengyao Shi, Haoda Zhang, Yiqi Cao, Yan Gu, Shen Li, Xiaoming Han, Sen Qian, and Shulin Liu. Study on the improvement of the 20-inch microchannel plate photomultiplier tubes for neutrino detector. *Nuclear Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers, Detectors and Associated Equipment*, 977:164333, 2020.

Full Authors List: LHAASO Collaboration

Zhen Cao^{1,2,3}, F. Aharonian^{4,5}, Q. An^{6,7}, Axikegu⁸, L.X. Bai⁹, Y.X. Bai^{1,3}, L.X. Bai⁹, Y.X. Bai^{1,3}, Y.W. Bao¹⁰, D. Bastieri¹¹, X.J. Bi^{1,2,3}, Y.J. Bi^{1,3}, H. Cai¹², J.T. Cai¹¹, Zhe Cao^{6,7}, J. Chang¹³, J.F. Chang^{1,3,6}, B.M. Chen¹⁴, E.S. Chen^{1,2,3}, J. Chen⁹, Liang Chen^{1,2,3}, Liang Chen¹⁵, Long Chen⁸, M.J. Chen^{1,3}, M.L. Chen^{1,3,6}, Q.H. Chen⁸, S.H. Chen^{1,2,3}, S.Z. Chen^{1,3}, T.L. Chen¹⁶, X.L. Chen^{1,2,3}, Y. Chen¹⁰, N. Cheng^{1,3}, Y.D. Cheng^{1,3}, S.W. Cui¹⁴, X.H. Cui¹⁷, Y.D. Cui¹⁸, B. D'Ettorre Piazzoli¹⁹, B.Z. Dai²⁰, H.L. Chen^{1,3,6}, Z.G. Dai⁷, N. Cheng^{1,5}, Y.D. Cheng^{1,5}, S.W. Cul^{1,7}, X.H. Cul^{1,7}, Y.D. Cul^{1,5}, B. D Ettorre Plazzoli^{1,5}, B.Z. Dal^{2,6}, H.L. Dai^{1,3,6}, Z.G. Dai⁷, Danzengluobu¹⁶, D. della Volpe²¹, X.J. Dong^{1,3}, K.K. Duan¹³, J.H. Fan¹¹, Y.Z. Fan¹³, Z.X. Fan^{1,3}, J. Fang²⁰, K. Fang^{1,3}, C.F. Feng²², L. Feng¹³, S.H. Feng^{1,3}, Y.L. Feng¹³, B. Gao^{1,3}, C.D. Gao²², L.Q. Gao^{1,2,3}, Q. Gao¹⁶, W. Gao²², M.M. Ge²⁰, L.S. Geng^{1,3}, G.H. Gong²³, Q.B. Gou^{1,3}, M.H. Gu^{1,3,6}, F.L. Guo¹⁵, J.G. Guo^{1,2,3}, X.L. Guo⁸, Y.Q. Guo^{1,3}, Y.Y. Guo^{1,2,3,13}, Y.A. Han²⁴, H.H. He^{1,2,3}, H.N. He¹³, J.C. He^{1,2,3}, S.L. He¹¹, X.B. He¹⁸, Y. He⁸, M. Heller²¹, Y.K. Hor¹⁸, C. Hou^{1,3}, H.B. Hu^{1,2,3}, S. Hu⁹, S.C. Hu^{1,2,3}, X.J. Hu²³, D.H. Huang⁸, Q.L. Huang^{1,3}, W.H. Huang²², X.T. Huang²², X.Y. Huang¹³, Z.C. Huang⁸, F. Ji^{1,3}, X.L. Ji^{1,3,6}, H.Y. Jia⁸, K. Jiang^{6,7}, Z.J. Jiang²⁰, C. Jin^{1,2,3}, T. Ke^{1,3}, D. Kuleshov²⁵, K. Levochkin²⁵, B.B. Li¹⁴, Cheng Li^{6,7}, Cong Li^{1,3}, F. Li^{1,3,6}, H.B. Li^{1,3}, H.C. Li^{1,3}, H.Y. Li^{7,13}, J. Li^{1,3,6}, K. Li^{1,3}, W.L. Li²², X.R. Li^{1,3}, Xin Li^{6,7}, Xin Li⁸, Y. Li⁹, Y.Z. Li^{1,2,3}, Zhe $\begin{array}{l} Li^{1,3}, Zhuo \ Li^{26}, E.W. \ Liang^{27}, Y.F. \ Liang^{27}, S.J. \ Lin^{18}, B. \ Liu^{7}, C. \ Liu^{1,3}, D. \ Liu^{22}, H. \ Liu^{8}, H.D. \ Liu^{24}, J. \ Liu^{1,3}, J.L. \ Liu^{28}, J.S. \ Liu^{18}, J.Y. \ Liu^{1,3}, M.Y. \ Liu^{16}, R.Y. \ Liu^{10}, S.M. \ Liu^{8}, W. \ Liu^{1,3}, Y. \ Liu^{11}, Y.N. \ Liu^{23}, Z.X. \ Liu^{9}, W.J. \ Long^{8}, R. \ Lu^{20}, H.K. \ Lv^{1,3}, B.Q. \ Ma^{26}, L.L. \ Ma^{1,3}, X.H. \ Ma^{1,3}, J.R. \ Mao^{29}, A. \ Mascod^{8}, Z. \ Min^{1,3}, W. \ Mitthumsiri^{30}, T. \ Montaruli^{21}, Y.C. \ Nan^{22}, B.Y. \ Pang^{8}, \end{array}$ P. Pattarakijwanich³⁰, Z.Y. Pei¹¹, M.Y. Qi^{1,3}, Y.Q. Qi¹⁴, B.Q. Qiao^{1,3}, J.J. Qin⁷, D. Ruffolo³⁰, V. Rulev²⁵, A. Sáiz³⁰, L. Shao¹⁴, O. Shchegolev^{25,31}, X.D. Sheng^{1,3}, J.Y. Shi^{1,3}, H.C. Song²⁶, Yu.V. Stenkin^{25,31}, V. Stepanov²⁵, Y. Su³², Q.N. Sun⁸, X.N. Sun²⁷, Z.B. Sun³³, P.H.T. Tam¹⁸, Z.B. Tang^{6,7}, W.W. Tian^{2,17}, B.D. Wang^{1,3}, C. Wang³³, H. Wang⁸, H.G. Wang¹¹, J.C. Wang²⁹, J.S. Wang²⁸, L.P. Wang²², L.Y. Wang^{1,3}, R.N. Wang⁸, W. Wang¹⁸, W. Wang¹², X.G. Wang²⁷, X.J. Wang^{1,3}, X.Y. Wang¹⁰, Y. Wang⁸, Y.D. Wang^{1,3}, Y.J. Wang^{1,2}, L.I. Wang^{1,2}, K.N. Wang², W. Wang^{2,2}, W. Wang^{2,2}, X.O. Wang^{2,2}, X.J. Wang^{3,2}, X.I. Wang^{1,2}, T. Wang^{2,1}, I.D. Wang^{2,3}, T.J. Wang^{1,2,3}, Z.H. Wang^{9,2}, Z.X. Wang²⁰, Zhen Wang²⁸, Zheng Wang^{1,3,6}, D.M. Wei¹³, J.J. Wei¹³, Y.J. Wei^{1,2,3}, T. Wen²⁰, C.Y. Wu^{1,3}, H.R. Wu^{1,3}, S. Wu^{1,3}, W.X. Wu⁸, X.F. Wu¹³, S.Q. Xi^{1,3}, J. Xia^{7,13}, J.J. Xia⁸, G.M. Xiang^{2,15}, D.X. Xiao¹⁶, G. Xiao^{1,3}, H.B. Xiao¹¹, G.G. Xin¹², Y.L. Xin⁸, Y. Xing¹⁵, D.L. Xu²⁸, R.X. Xu²⁶, L. Xue²², D.H. Yan²⁹, J.Z. Yan¹³, C.W. Yang⁹, F.F. Yang^{1,3,6}, J.Y. Yang¹⁸, L.L. Yang¹⁸, M.J. Yang^{1,3}, R.Z. Yang⁷, S.B. Yang²⁰, Y.H. Yao⁹, Z.G. Yao^{1,3}, Y.M. Ye²³, L.Q. Yin^{1,3}, N. Yin²², X.H. You^{1,3}, Z.Y. You^{1,2,3}, Y.H. Yu²², Q. Yuan¹³, H.D. Zeng¹³, T.X. Zeng^{1,3,6}, W. Zeng²⁰, Z.K. Zeng^{1,2,3}, M. Zha^{1,3}, X.X. Zhai^{1,3}, B.B. Zhang¹⁰, H.M. Zhang¹⁰, H.Y. Zhang²², J.L. Zhang¹⁷, J.W. Zhang⁹, L.X. Zhang¹¹, Li Zhang²⁰, Lu Zhang¹⁴, P.F. Zhang²⁰, P.P. Zhang¹⁴, R. Zhang^{7,13}, S.R. Zhang¹⁴, S.S. Zhang^{1,3}, X. Zhang¹⁰, X.P. Zhang^{1,3}, Y.F. Zhang⁸, Y.L. Zhang^{1,3}, Yi Zhang^{1,13}, Yong Zhang^{1,3}, B. Zhao⁸, J. Zhao^{1,3}, L. Zhao^{6,7}, L.Z. Zhao¹⁴, S.P. Zhao^{13,22}, F. Zheng³³, Y. Zheng⁸, B. Zhou^{1,3}, H. Zhou²⁸, J.N. Zhou¹⁵, P. Zhou¹⁰, R. Zhou⁹, X.X. Zhou⁸, C.G. Zhu²², F.R. Zhu⁸, H. Zhu¹⁷, K.J. Zhu^{1,2,3,6} and X. Zuo^{1,3}

¹Key Laboratory of Particle Astrophyics & Experimental Physics Division & Computing Center, Institute of High Energy Physics, Chinese Academy of Sciences, 100049 Beijing, China.

²University of Chinese Academy of Sciences, 100049 Beijing, China.

³TIANFU Cosmic Ray Research Center, Chengdu, Sichuan, China.

⁴Dublin Institute for Advanced Studies, 31 Fitzwilliam Place, 2 Dublin, Ireland.

⁵Max-Planck-Institut for Nuclear Physics, P.O. Box 103980, 69029 Heidelberg, Germany.

⁶State Key Laboratory of Particle Detection and Electronics, China.

⁷University of Science and Technology of China, 230026 Hefei, Anhui, China.

⁸School of Physical Science and Technology & School of Information Science and Technology, Southwest Jiaotong University, 610031 Chengdu, Sichuan, China.

⁹College of Physics, Sichuan University, 610065 Chengdu, Sichuan, China.

¹⁰School of Astronomy and Space Science, Nanjing University, 210023 Nanjing, Jiangsu, China.

¹¹Center for Astrophysics, Guangzhou University, 510006 Guangzhou, Guangdong, China.

¹²School of Physics and Technology, Wuhan University, 430072 Wuhan, Hubei, China.

¹³Key Laboratory of Dark Matter and Space Astronomy, Purple Mountain Observatory, Chinese Academy of Sciences, 210023 Nanjing, Jiangsu, China.

¹⁴Hebei Normal University, 050024 Shijiazhuang, Hebei, China.

¹⁵Key Laboratory for Research in Galaxies and Cosmology, Shanghai Astronomical Observatory, Chinese Academy of Sciences, 200030 Shanghai, China.

¹⁶Key Laboratory of Cosmic Rays (Tibet University), Ministry of Education, 850000 Lhasa, Tibet, China.

¹⁷National Astronomical Observatories, Chinese Academy of Sciences, 100101 Beijing, China.

¹⁸School of Physics and Astronomy & School of Physics (Guangzhou), Sun Yat-sen University, 519000 Zhuhai, Guangdong,

China.

¹⁹Dipartimento di Fisica dell'Università di Napoli [`]Federico II', Complesso Universitario di Monte Sant'Angelo, via Cinthia, 80126 Napoli, Italy.

²⁰School of Physics and Astronomy, Yunnan University, 650091 Kunming, Yunnan, China.

²¹Departement de Physique Nucleaire et Corpusculaire, Faculte de Sciences, Universite de Gen^eve, 24 Quai Ernest Ansermet, 1211 Geneva, Switzerland.

²²Institute of Frontier and Interdisciplinary Science, Shandong University, 266237 Qingdao, Shandong, China.

²³Department of Engineering Physics, Tsinghua University, 100084 Beijing, China.

²⁴School of Physics and Microelectronics, Zhengzhou University, 450001 Zhengzhou, Henan, China.

²⁵Institute for Nuclear Research of Russian Academy of Sciences, 117312 Moscow, Russia.

²⁶School of Physics, Peking University, 100871 Beijing, China.

²⁷School of Physical Science and Technology, Guangxi University, 530004 Nanning, Guangxi, China.

²⁸Tsung-Dao Lee Institute & School of Physics and Astronomy, Shanghai Jiao Tong University, 200240 Shanghai, China.

²⁹Yunnan Observatories, Chinese Academy of Sciences, 650216 Kunming, Yunnan, China.

³⁰Department of Physics, Faculty of Science, Mahidol University, 10400 Bangkok, Thailand.

³¹Moscow Institute of Physics and Technology, 141700 Moscow, Russia.

³²Key Laboratory of Radio Astronomy, Purple Mountain Observatory, Chinese Academy of Sciences, 210023 Nanjing, Jiangsu, China.

³³National Space Science Center, Chinese Academy of Sciences, 100190 Beijing, China.