

Effects of Residue Background Events in Direct Detection Experiments on Determining Properties of Halo Dark Matter

Chung-Lin Shan*

Department of Physics, National Cheng Kung University, Tainan City 70101, Taiwan, R.O.C.

Physics Division, National Center for Theoretical Sciences, Hsinchu City 30013, Taiwan, R.O.C.

E-mail: clshan@mail.ncku.edu.tw

We reexamine the model-independent data analysis methods for extracting properties of Weakly Interacting Massive Particles (WIMPs) by using data (measured recoil energies) from direct Dark Matter detection experiments directly and, as a more realistic study, consider a small fraction of residue background events, which pass all discrimination criteria and then mix with other real WIMP-induced signals in the analyzed data sets. In this talk, the effects of residue backgrounds on the determination of the mass of halo Dark Matter particle as well as on the reconstruction of its one-dimensional velocity distribution function will be discussed.

Identification of dark matter 2010

July 26-30, 2010

Montpellier, France

*Speaker.

1. Introduction

In our earlier work on the development of model-independent data analysis methods for extracting properties of Weakly Interacting Massive Particles (WIMPs) by using measured recoil energies from direct Dark Matter detection experiments directly [1–3], it was assumed that the analyzed data sets are background-free, i.e., all events are WIMP signals. Active background discrimination techniques should make this condition possible. For example, by using the ratio of the ionization to recoil energy, the so-called “ionization yield”, combined with the “phonon pulse timing parameter”, the CDMS-II collaboration claimed that electron recoil events can be rejected event-by-event with a misidentification fraction of $< 10^{-6}$. [4] The CRESST collaboration demonstrated also that the pulse shape discrimination (PSD) technique can distinguish WIMP-induced nuclear recoils from those induced by backgrounds by means of inserting a scintillating foil, which causes some additional scintillation light for events induced by α -decay of ^{210}Po and thus shifts the pulse shapes of these events faster than pulses induced by WIMP interactions in the crystal [5].¹

However, as the most important issue in all underground experiments, possible residue background events which pass all discrimination criteria and then mix with other real WIMP-induced events in our data sets should also be considered. Therefore, as a more realistic study, we take into account small fractions of residue background events mixed in experimental data sets and want to study how well the model-independent methods could extract the *input* WIMP properties by using these “impure” data sets and how “dirty” these data sets could be to be still useful.

¹More details about background discrimination techniques and status see also e.g., Refs. [6].

In this article, I focus on two properties of halo WIMPs: their mass m_χ and one-dimensional velocity distribution function $f_1(v)$. More detailed discussions can be found in Refs. [7, 8].

2. Effects of residue background events

In our numerical simulations based on the Monte Carlo method, while the shifted Maxwellian velocity distribution [9, 1] with the standard values of the Sun’s orbital velocity and the Earth’s velocity in the Galactic frame: $v_0 \simeq 220$ km/s and $v_e = 1.05 v_0$, and the Woods-Saxon form for the elastic nuclear form factor for the spin-independent (SI) WIMP-nucleus interaction [10, 9] have been used for generating WIMP-induced signals, a *target-dependent exponential* form for residue background events has been introduced [7]:

$$\left(\frac{dR}{dQ}\right)_{\text{bg,ex}} = \exp\left(-\frac{Q/\text{keV}}{A^{0.6}}\right). \quad (2.1)$$

Here Q is the recoil energy, A is the atomic mass number of the target nucleus. The power index of A , 0.6, is an empirical constant, which has been chosen so that the exponential background spectrum is somehow *similar to*, but still *different from* the expected recoil spectrum of the target nucleus (see Figs. 1); otherwise, there is in practice no difference between the WIMP scattering and background spectra. Note that, the atomic mass number A has been used here just as the simplest, unique characteristic parameter in the analytic form (2.1) for defining the residue background spectrum for *different* target nuclei. It does *not* mean that the (superposition of the real) background spectra would depend simply/primarily on A or on the mass of the target nucleus, m_N .

Note also that, firstly, the exponential form (2.1) for residue background spectrum is rather naive; however, since we consider here *only a few (tens) residue* background events induced

by perhaps *two or more* different sources, pass all discrimination criteria, and then mix with other WIMP–induced events in our data sets of $\mathcal{O}(100)$ total events, exact forms of different background spectra are actually not very important and this exponential spectrum should practically not be unrealistic. Secondly, our model–independent data analysis procedures requires only measured recoil energies from one or more experimental data sets with different target nuclei [1–3]. Hence, for applying these methods to future real direct detection data, the prior knowledge about (different) background source(s) is *not required at all*.

Moreover, the maximal cut–off of the velocity distribution function has been set as $v_{\max} = 700$ km/s. The experimental threshold energy has been assumed to be negligible and the maximal cut–off energy is set as 100 keV. The background window (the possible energy range in which residue background events *cannot be ignored*, compared to some other ranges) has been assumed to be the same as the experimental possible energy range. Note here that the actual numbers of generated signal and background events in each simulated experiment are Poisson–distributed around their expectation values *independently*, and the total event number in one experiment is then the sum of these two numbers; both generated signal and background events are treated as WIMP signals in our analyses. Additionally, we assumed that all experimental systematic uncertainties as well as the uncertainty on the measurement of the recoil energy could be ignored.

2.1 On the measured recoil spectrum

In Figs. 1 I show measured energy spectra (solid red histograms) for a ^{76}Ge target with three different WIMP masses: 25 (top), 50 (middle), and 250 (bottom) GeV. The dotted blue curves show the elastic WIMP–nucleus scattering spectra, the dashed green curves indicate

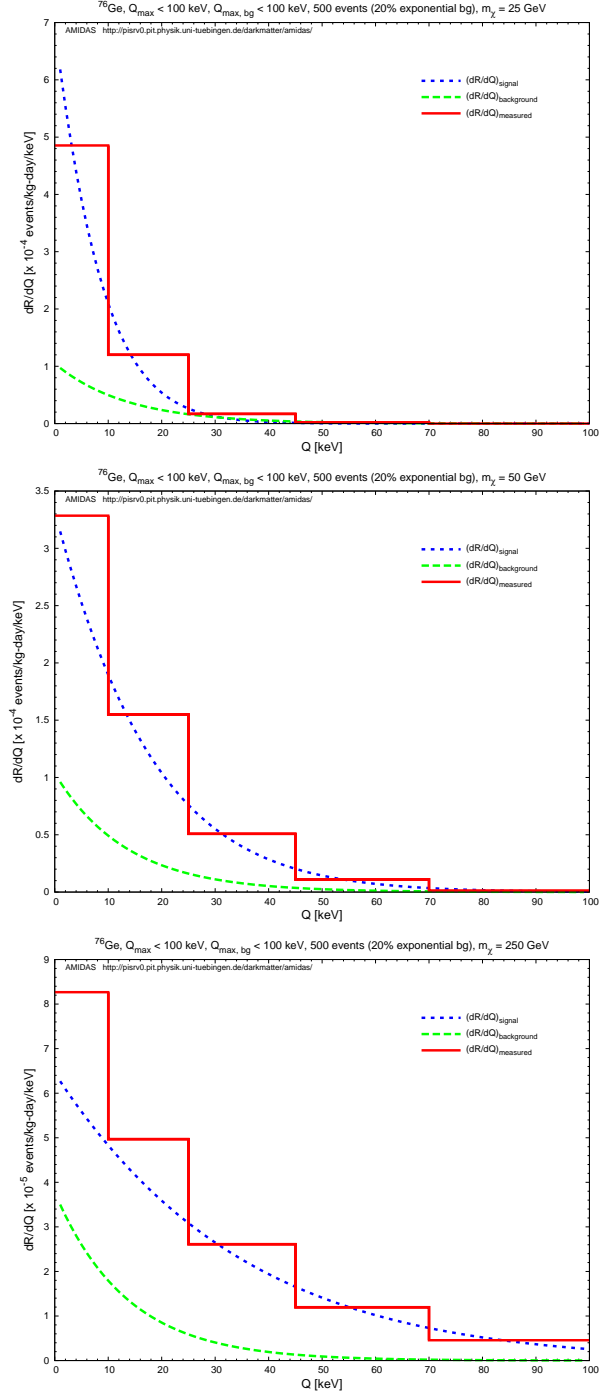


Figure 1: Measured energy spectra (solid red histograms) for a ^{76}Ge target with three different WIMP masses: 25 (top), 50 (middle), and 250 (bottom) GeV. The dotted blue curves are the elastic WIMP–nucleus scattering spectra, whereas the dashed green curves are the exponential background spectra normalized to fit to the chosen background ratio, which has been set as 20% here (plots from Ref. [7]).

the exponential background spectrum given in Eq. (2.1), which have been normalized so that the ratios of the areas under these background spectra to those under the (dotted blue) WIMP scattering spectra are equal to the background–signal ratio in the whole data sets. 5,000 experiments with 500 total events on average in each experiment have been simulated.

It can be found here that, the shape of the WIMP scattering spectrum depends highly on the WIMP mass: for light WIMPs ($m_\chi \lesssim 50$ GeV), the recoil spectra drop sharply with increasing recoil energies, while for heavy WIMPs ($m_\chi \gtrsim 100$ GeV), the spectra become flatter. In contrast, the exponential background spectra shown here depend only on the target mass and are rather *flatter/sharper* for *light/heavy* WIMP masses compared to the WIMP scattering spectra. This means that, once input WIMPs are *light/heavy*, background events would contribute relatively more to *high/low* energy ranges, and, consequently, the measured energy spectra would mimic scattering spectra induced by *heavier/lighter* WIMPs.

2.2 On determining the WIMP mass

Fig. 2 show the *median* values of the reconstructed WIMP mass and the lower and upper bounds of the 1σ statistical uncertainty by means of the model–independent procedure introduced in Refs. [2] with mixed data sets from WIMP–induced and background events as functions of the input WIMP mass. As in Refs. [2], ^{28}Si and ^{76}Ge have been chosen as two target nuclei. The background ratios shown here are no background (dashed green), 10% (long–dotted blue), 20% (solid red), and 40% (dash–dotted cyan) background events in the analyzed data sets. $2 \times 5,000$ experiments with 50 total events on average in each experiment have been simulated.

It can be seen here clearly that, since for *light* WIMP masses ($m_\chi \lesssim 100$ GeV), due to

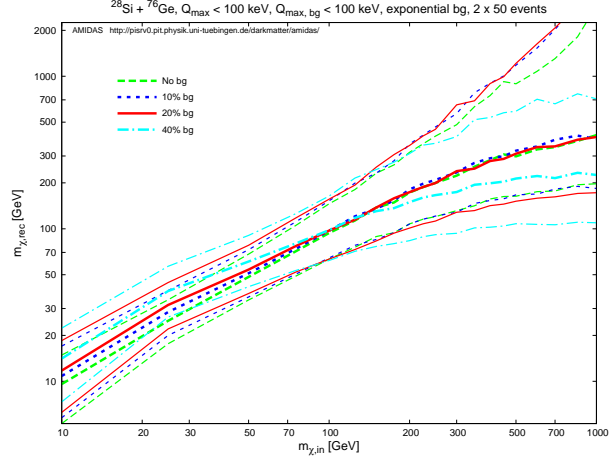


Figure 2: The reconstructed WIMP masses as functions of the input WIMP mass. ^{28}Si and ^{76}Ge have been chosen as two target nuclei. The background ratios shown here are no background (dashed green), 10% (long–dotted blue), 20% (solid red), and 40% (dash–dotted cyan) background events in the analyzed data sets. Each experiment contains 50 total events on average. Other parameters are as in Figs. 1 (plot from Ref. [7]).

the relatively flatter background spectrum (compared to the scattering spectrum induced by WIMPs) or, in practice, some background sources in high energy ranges, the energy spectrum of all recorded events would mimic a scattering spectrum induced by WIMPs with a relatively *heavier* mass, the reconstructed WIMP masses as well as the statistical uncertainty intervals could be *overestimated*. In contrast, for *heavy* WIMP masses ($m_\chi \gtrsim 100$ GeV), due to the relatively sharper background spectrum or e.g., some electronic noise, relatively more background events contribute to low energy ranges, the energy spectrum of all recorded events would thus mimic a scattering spectrum induced by WIMPs with a relatively *lighter* mass. Hence, the reconstructed WIMP masses as well as the statistical uncertainty intervals could be *underestimated*. Nevertheless, Fig. 2 shows that, with $\sim 20\%$ residue background events in the analyzed data sets of ~ 50 total events, the 1σ statistical uncertainty band can cover the true

WIMP mass pretty well; if WIMPs are light ($m_\chi \lesssim 200$ GeV), the maximal acceptable fraction of residue background events could even be as large as $\sim 40\%$.

2.3 On reconstructing the one-dimensional WIMP velocity distribution function

In this section I show the *median* values of the reconstructed one-dimensional velocity distribution function of halo WIMPs² with its 1σ statistical uncertainty by means of the model-independent method introduced in Ref. [1] with mixed data sets. As in Ref. [1], a ^{76}Ge nucleus has been chosen as our detector target for reconstructing $f_1(v)$; while a ^{28}Si target and a *second* ^{76}Ge target have been used for determining m_χ . The background ratios shown here are no background (dashed green), 10% (long-dotted blue), and 20% (solid red) background events in the analyzed data set(s). ($3 \times$) 5,000 experiments with 500 total events on average in each experiment have been simulated.

2.3.1 With a precisely known WIMP mass

In Figs. 3 we first assume that the required WIMP mass for reconstructing $f_1(v)$ has been known precisely with a negligible uncertainty. The horizontal bars show the sizes of the windows used for estimating $f_1(v)$ [1]. Five bins have been used and up to three bins have been combined to a window, in order to collect more events in wider ranges and thereby reduce the statistical uncertainty on the reconstructed $f_1(v)$.³

It can be seen that, since for *heavy* WIMP masses ($m_\chi \gtrsim 100$ GeV), the relatively sharper background spectrum contributes more events to low energy ranges, or, equivalently, to low

²Note that, since the experimental maximal cut-off energy is fixed as 100 keV, for heavy input WIMP masses ($m_\chi \gtrsim 250$ GeV), one can reconstruct the velocity distribution function only in the velocity range $v \lesssim 300$ km/s.

³Note that, since the neighboring windows overlap, the estimates of $f_1(v)$ at adjacent points are correlated.

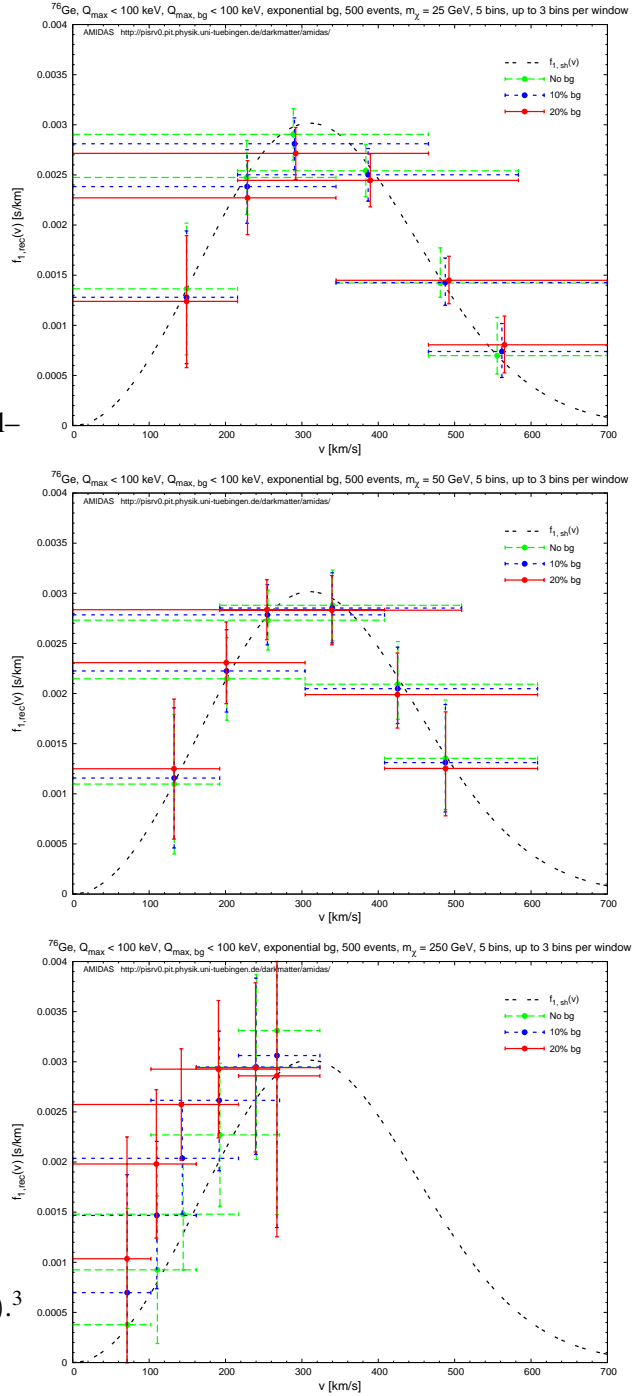


Figure 3: The reconstructed one-dimensional WIMP velocity distribution function for three different WIMP masses: 25 (top), 50 (middle), and 250 (bottom) GeV. The double-dotted black curves are the input shifted Maxwellian velocity distribution. The background ratios shown here are no background (dashed green), 10% (long-dotted blue), and 20% (solid red) background events in the analyzed data set. Each experiment contains 500 total events on average. Other parameters are as in 5 Figs. 1 (plots from Ref. [8]).

velocity ranges, the reconstructed velocity distribution would be shifted to *lower* velocities. For an input WIMP mass of 100 GeV and the background ratio of 10% (20%), the peak of the reconstructed velocity distribution function could be shifted by ~ 30 (60) km/s. In contrast, for *light* WIMP masses ($m_\chi \lesssim 50$ GeV), the relatively flatter background spectrum contributes more events to high energy/velocity ranges, and the reconstructed velocity distribution would be shifted to *higher* velocities. Moreover, our simulation results indicate that, with an $\sim 10\% - 20\%$ background ratio in the analyzed data set of ~ 500 total events, one could in principle reconstruct the one-dimensional velocity distribution function of halo WIMPs with an $\sim -6.5\%$ (for a 25 GeV WIMP mass, 20% background events) $- \sim +38\%$ (for a 250 GeV WIMP mass, 10% background events) deviation. If the mass of halo WIMPs is $\mathcal{O}(50$ GeV), the maximal acceptable background ratio could even be as large as $\sim 40\%$ with a deviation of only $\sim +14\%$.

2.3.2 With a reconstructed WIMP mass

In Figs. 4 the required WIMP mass for reconstructing $f_1(v)$ has been reconstructed with *other* direct detection experiments. Note that, while the vertical bars show the 1σ statistical uncertainties on the reconstructed $f_1(v)$, taken into account the statistical uncertainty on the reconstructed WIMP mass, the horizontal bars indicate here the 1σ statistical uncertainties on the estimates of the reconstructed v points, due to the uncertainty on the reconstructed m_χ .

It has been found that, firstly, as shown in Figs. 2, for an input WIMP mass of 100 GeV, the reconstructed mass doesn't differ very much from the true value. Hence, the reconstructed $f_1(v)$ is approximately the same for both cases with the input and reconstructed WIMP masses. However, for *light/heavy* input masses ($m_\chi \lesssim / \gtrsim 100$ GeV), the reconstructed $f_1(v)$ with the

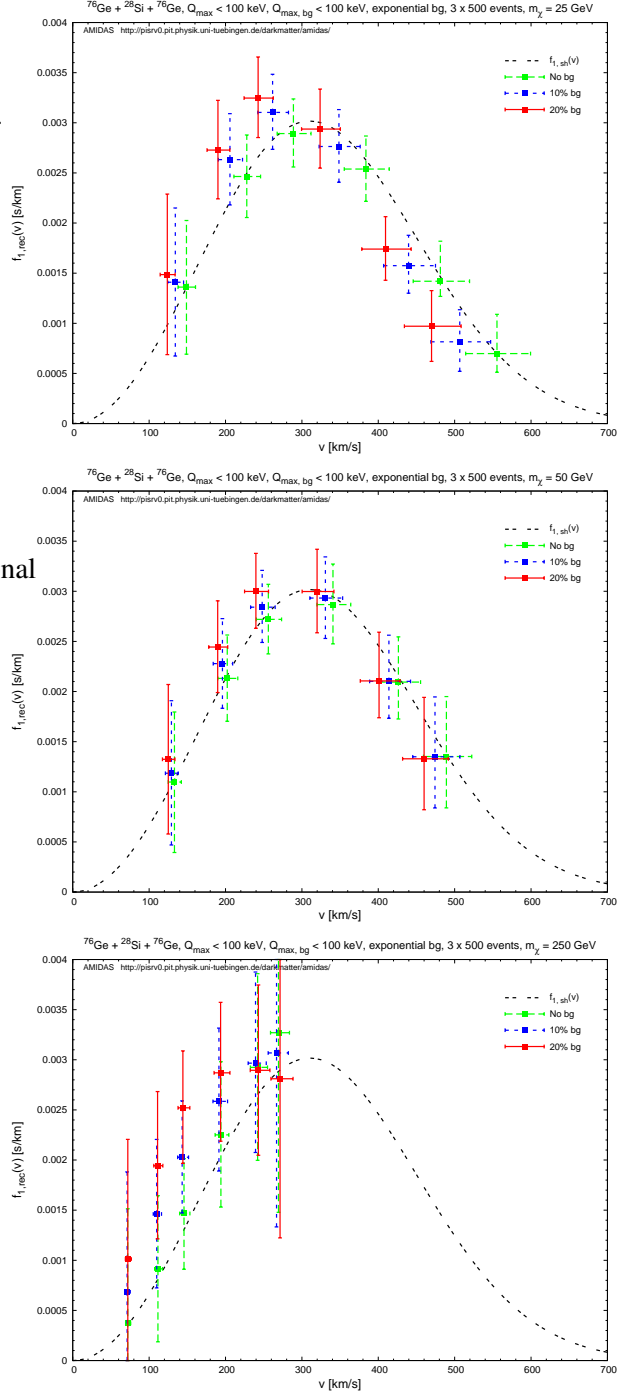


Figure 4: As in Figs. 3, except that the WIMP masses have been reconstructed by means of the procedure introduced in Refs. [2] (plots from Ref. [8]).

reconstructed WIMP mass shift to relatively *lower/higher* velocities compared to the case with the input (true) WIMP mass. This effect

caused directly by the over-/underestimate of the reconstructed WIMP mass [8]: once the reconstructed WIMP mass is over-/underestimated from the real value, a transformation constant from Q to v will thus be under-/overestimated. Consequently, the reconstructed v points will be smaller/larger than the true values.

In contrast to the first effect discussed in the previous subsection, this second effect draws the reconstructed WIMP velocity distribution (more strongly) to the opposite directions. Nevertheless, with an $\sim 5\% - 10\%$ background ratio in the analyzed data sets of ~ 500 total events, one could in principle still reconstruct the one-dimensional WIMP velocity distribution function with an $\sim +7\%$ (for 25 GeV WIMPs, 10% backgrounds) $- \sim +16\%$ (for 250 GeV WIMPs, 5% backgrounds) deviation. If $m_\chi \lesssim 100$ GeV, the maximal acceptable background ratio could even be as large as $\sim 20\%$ with a deviation of only $\sim +9\%$.

3. Summary

In this article we reexamine the data analysis methods introduced in Refs. [1, 2] for determining the mass and one-dimensional velocity distribution function of halo Dark Matter particle from measured recoil energies of direct detection experiments directly, by taking into account small fractions of residue background events, which pass all discrimination criteria and then mix with other real WIMP-induced events in the analyzed data sets.

Our simulations show that, with a background ratio of $\sim 5\% - 10\%$ in data sets of ~ 500 total events, the one-dimensional WIMP velocity distribution function can in principle be reconstructed with an $\sim -5\% - +20\%$ systematic deviation; whereas with a background ratio of even $\sim 10\% - 20\%$ in data sets of only ~ 50 total events, the 1σ statistical uncertainty

band of the reconstructed WIMP mass can cover the true value pretty well.

Acknowledgments

The author would like to thank the Physikalisches Institut der Universität Tübingen for the technical support of the computational work demonstrated in this article. This work was partially supported by the National Science Council of R.O.C. under contract no. NSC-98-2811-M-006-044 as well as by the Focus Group on Cosmology and Particle Astrophysics, National Center of Theoretical Sciences, R.O.C..

References

- [1] M. Drees and C.-L. Shan, *JCAP* **0706**, 011 (2007).
- [2] M. Drees and C.-L. Shan, [arXiv:0710.4296](https://arxiv.org/abs/0710.4296) (2007); *JCAP* **0806**, 012 (2008).
- [3] M. Drees and C.-L. Shan, *PoS IDM2008*, 110 (2008); [arXiv:0903.3300](https://arxiv.org/abs/0903.3300) (2009).
- [4] CDMS Collab., Z. Ahmed *et al.*, *Science* **327**, 1619 (2010).
- [5] CRESST Collab., R. F. Lang *et al.*, *Astropart. Phys.* **33**, 60 (2010); CRESST Collab., J. Schmaler *et al.*, *AIP Conf. Proc.* **1185**, 631 (2009).
- [6] E. Aprile and L. Baudis, for the XENON100 Collab., *PoS IDM2008*, 018 (2008); EDELWEISS Collab., A. Broniatowski *et al.*, *Phys. Lett.* **B 681**, 305 (2009); EDELWEISS Collab., E. Armengaud *et al.*, *Phys. Lett.* **B 687**, 294 (2010); CRESST Collab., R. F. Lang *et al.*, *Astropart. Phys.* **32**, 318 (2010).
- [7] Y.-T. Chou and C.-L. Shan, *JCAP* **1008**, 014 (2010).
- [8] C.-L. Shan, *JCAP* **1006**, 029 (2010).
- [9] G. Jungman, M. Kamionkowski and K. Griest, *Phys. Rep.* **267**, 195 (1996).
- [10] J. Engel, *Phys. Lett.* **B 264**, 114 (1991).