

# Constraints and prospects on gravitational wave and neutrino emission using GW150914

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The recent LIGO observation of gravitational waves from a binary black hole merger triggered several follow-up searches from both electromagnetic wave as well as neutrino observatories. Since in general, it is expected that all matter has been removed from the binary black hole environment long before the merger, no neutrino emission is expected from such mergers. Still, it remains interesting to test this hypothesis. The ratio of the energy emitted in neutrinos with respect to gravitational waves represents a useful parameter to constrain the environment of such astrophysical events. In addition to putting constraints by use of the non-detection of counterpart neutrinos, it is also possible to consider the diffuse neutrino flux measured by the IceCube observatory as the maximum contribution from an extrapolated full class of BBHs. Both methods currently lead to similar bounds on the fraction of energy that can be emitted in neutrinos. Nevertheless, combining both methods should allow to strongly constrain the source population in case of a future neutrino counterpart detection. The proposed approach can and will be applied to potential upcoming LIGO events, including binary neutron stars and black hole-neutron star mergers, for which a neutrino counterpart is expected.

*35th International Cosmic Ray Conference*

*10-20 July, 2017*

*Bexco, Busan, Korea*

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## 1. Introduction

With the first direct detection of gravitational waves, a new observational window on the universe has opened. Following the multimessenger paradigm, gravitational waves can provide additional information about known astrophysical objects or reveal those that are otherwise unobservable.

The first detected event, GW150914 [1], has a signal compatible with the merger of a black hole binary. In this event, two black holes, with masses equal to  $36^{+5}_{-4} M_{\odot}$  and  $29^{+4}_{-4} M_{\odot}$ , merged to form a single black hole of mass  $62^{+4}_{-4} M_{\odot}$ , releasing  $3^{+0.5}_{-0.5} M_{\odot}$  of energy into gravitational waves. A black hole binary system is expected to be free of all matter by the time of the merger [2]. Therefore, no emission is expected from such a system except for gravitational waves. Now that a binary black hole (BBH) merger has been detected, this expectation can be tested.

The detection of GW150914 triggered multiple follow-up searches for electromagnetic [3] and neutrino [4, 5, 6] signals. While none of these searches detected a significant signal, Fermi GBM did see a sub-threshold event resembling a short gamma-ray burst (GRB) [7]. However, INTEGRAL imaged the same area and found no such signal [8], casting doubt on the interpretation of this observation. Still, several models were constructed that could explain a possible short GRB from BBH mergers (e.g. [2, 9, 10]).

Given that Megaton-scale neutrino detectors such as IceCube [11] are available, it is appropriate to test whether BBH mergers emit neutrinos. In the following, we will use the non-detection of signal neutrinos associated with GW150914 [4] to put a generic bound on the high energy neutrino emission from BBH mergers, relative to the energy emitted in gravitational waves [12]. This is compared to the current bound coming from the diffuse astrophysical neutrino flux detected by IceCube [13]. Afterwards, we show how to interpret this bound in terms of specific models. Using this interpretation, we estimate the amount of neutrino emission that could be expected in realistic models for BBH mergers. Conversely, from the current limit on the neutrino emission, we obtain a bound on the amount of matter present in the environment of BBH mergers.

## 2. Neutrino emission

The neutrino emission coming from (a class of) BBH mergers will be parametrised by the neutrino emission fraction

$$f_{\text{BBH}}^{\nu} = \frac{E_{\nu}}{E_{\text{GW}}}, \quad (2.1)$$

where  $E_{\nu}$  is the total amount of energy emitted in neutrinos, while  $E_{\text{GW}}$  is the total energy released in gravitational waves, given by LIGO. This  $f_{\text{BBH}}^{\nu}$  can be probed directly by experiments searching for neutrino emission coincident with the merger event.

In addition,  $f_{\text{BBH}}^{\nu}$  is constrained by the diffuse high energy astrophysical neutrino flux detected by IceCube [13], since BBH mergers have been happening throughout the history of the universe. We calculate this bound following the approach in [14, 15], assuming that BBH mergers emit neutrinos following an  $E^{-2}$ -spectrum. The diffuse neutrino flux due to (a class of) BBH mergers occurring with a rate  $R$  (in  $\text{Gpc}^{-3} \text{yr}^{-1}$ ) in the local universe is given by [16],

$$E^2 \frac{dN_{\nu}}{dE_{\nu}} \Big|_{\text{obs}} = \left( f_{\text{BBH}}^{\nu} t_H \frac{c}{4\pi} \xi_z \right) E^2 \frac{d\dot{N}_{\nu}}{dE_{\nu}} \Big|_{\text{inj}, f_{\text{BBH}}^{\nu}=1}, \quad (2.2)$$

with the total injected flux

$$E^2 \left. \frac{d\dot{N}_\nu}{dE_\nu} \right|_{\text{inj}, f_{\text{BBH}}^V=1} = R E^2 \phi(E_\nu). \quad (2.3)$$

Here,  $\phi(E_\nu)$  is the neutrino flux from a single BBH merger in case that  $f_{\text{BBH}}^V = 1$ , where the integrated neutrino energy is equal to  $E_{\text{GW}}$ . The cosmic evolution of the sources is contained in  $\xi_z$ . For a merger rate following the star formation rate (SFR) [17, 18], one has  $\xi_z \approx 2.4$ .

The rate of BBH mergers in the local universe is determined by LIGO. When accounting for different possible mass distributions, the 90% credible interval for this rate is given by [19]

$$R = 9 - 240 \text{ Gpc}^{-3} \text{ yr}^{-1}. \quad (2.4)$$

For BBH mergers with black holes similar to GW150914, the corresponding rate is given by [19],

$$R_{\text{GW150914}} = 3.4_{-2.8}^{+8.8} \text{ Gpc}^{-3} \text{ yr}^{-1}. \quad (2.5)$$

The diffuse neutrino emission from BBH mergers (Eq. 2.2) is constrained to be below the diffuse astrophysical neutrino flux measured by IceCube [20]

$$E^2 \Phi(E) = 0.84 \pm 0.3 \times 10^{-8} \text{ GeV cm}^{-2} \text{ s}^{-1} \text{ sr}^{-1}. \quad (2.6)$$

Therefore, by combining information from neutrino and gravitational wave observatories, the astrophysical bound constrains the source evolution (through  $\xi_z$ ) and environment (through  $f_{\text{BBH}}^V$ ).

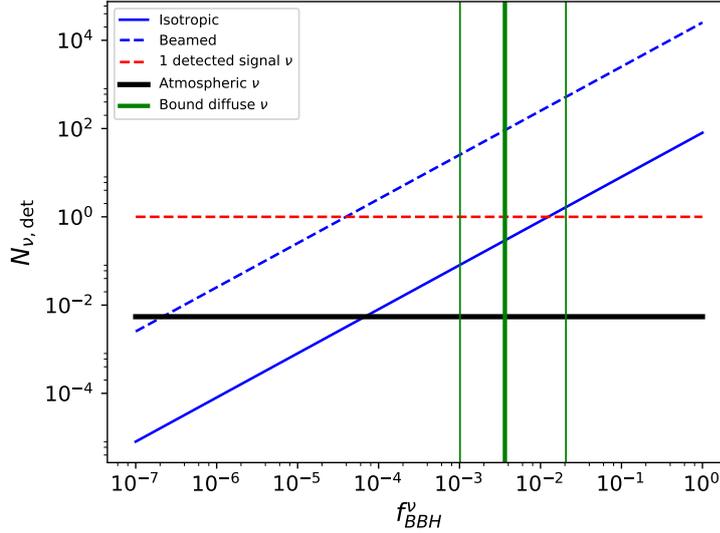
### 3. Limits from GW150914

Using the null detection of counterpart neutrinos to GW150914 by IceCube [4], we calculate the bound on the neutrino emission fraction  $f_{\text{BBH}}^V$ . Afterwards, we compare this with the constraint from the diffuse astrophysical neutrino flux for this class of mergers.

In order to find the constraint on  $f_{\text{BBH}}^V$  from the direct search, we calculate the expected neutrino flux that would be detectable on Earth, coming from GW150914, as a function of  $f_{\text{BBH}}^V$ . We assume that the neutrinos are emitted with an  $E^{-2}$ -spectrum from a distance of 410 Mpc (corresponding to the central value of the LIGO estimated distance of  $410_{-180}^{+160}$  Mpc). In order to convert the neutrino flux at Earth to the flux detected by IceCube, we use the effective area corresponding to the follow-up search [21] in the declination band  $-30^\circ < \delta < 0^\circ$ . We include the irreducible background from atmospheric neutrinos [22]. This background is integrated over a time interval of 1000 s, so that the signal is contained, and over a sky patch of  $600 \text{ deg}^2$ , which corresponds to the localization of GW150914 by LIGO.

Fig. 1 shows the expected number of detected neutrinos as a function of  $f_{\text{BBH}}^V$ , for both isotropic emission (full blue line) and beamed emission within a solid angle of  $0.2 \times 0.2$  (dashed blue line). From the non-detection of signal neutrinos,  $f_{\text{BBH}}^V$  is constrained to be below the level where one signal neutrino would be detected (red dashed line). This corresponds to  $f_{\text{BBH}}^V = 1.24 \times 10^{-2}$  for isotropic emission and  $f_{\text{BBH}}^V = 3.96 \times 10^{-5}$  for beamed emission. The irreducible background of atmospheric neutrinos is indicated by the solid black line. For a single BBH merger,

this background is far below the level where one signal neutrino could be detected and is thus negligible. In order to calculate the bound from the diffuse astrophysical neutrino flux, we only considered the class of BBH mergers similar to GW150914, occurring with the rate in Eq. 2.5. This bound is given by  $f_{\text{BBH}}^{\nu} = 1.01 \times 10^{-3} - 2.06 \times 10^{-2}$ , with a central value at  $f_{\text{BBH}}^{\nu} = 3.63 \times 10^{-3}$  (green lines).



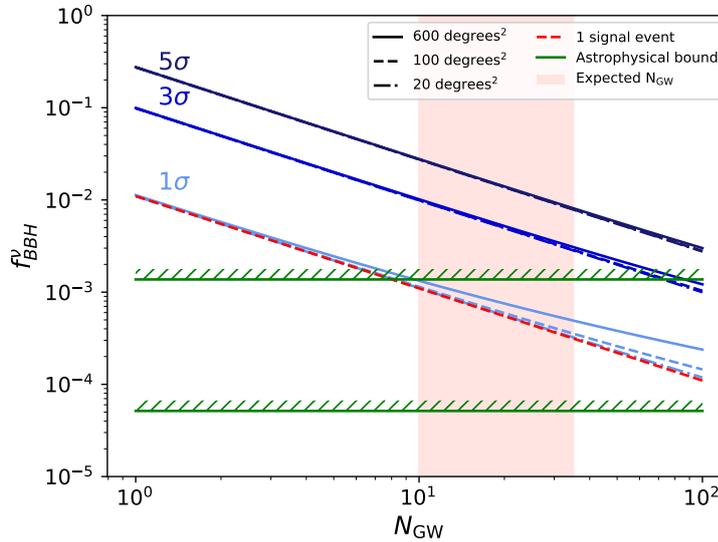
**Figure 1:** The expected number of neutrino events from GW150914, assuming an  $E^{-2}$  neutrino spectrum, as a function of  $f_{\text{BBH}}^{\nu}$ , for isotropic (full blue) and beamed (dashed blue) emission. The background, integrated over 1000 s and a solid angle of  $600 \text{ deg}^2$ , is shown by the black line. The red dashed line indicates 1 detected neutrino. The limits on  $f_{\text{BBH}}^{\nu}$  follows from the intersection of this line with the blue lines. The central value and uncertainty of the constraint from the astrophysical neutrino flux are shown in green. Figure from [12].

#### 4. Prospects

As more BBH mergers are detected, it will be possible to improve the limit on the neutrino emission by performing a stacked analysis. It will be assumed that all BBH mergers are similar to GW150914, injecting  $3 M_{\odot}$  of energy in gravitational waves and  $f_{\text{BBH}}^{\nu} \times 3 M_{\odot}$  of energy into neutrinos from a distance of 410 Mpc. Only the case of isotropic emission will be shown, while the results for beamed emission can be immediately obtained by rescaling with the solid angle  $\frac{4\pi}{\Delta\Omega} = \frac{4\pi}{0.2 \times 0.2}$ . The analysis from the previous section is repeated, with an effective area averaged over the full sky. Since the localization of GW events is expected to improve in future observation runs [23], the bound on  $f_{\text{BBH}}^{\nu}$  is calculated for background integrated over sky patches of  $600 \text{ deg}^2$ ,  $100 \text{ deg}^2$  and  $20 \text{ deg}^2$ . Because we now consider multiple events, the astrophysical bound is calculated with the full rate estimate in Eq. 2.4.

Fig. 2 shows the expected limit on  $f_{\text{BBH}}^{\nu}$  as a function of the number of detected BBH mergers ( $N_{\text{GW}}$ ) by LIGO. The limit on  $f_{\text{BBH}}^{\nu}$  from the stacked search with a signal strength  $S/\sqrt{S+B} = 5, 3, 1$  is given by the blue lines (from top to bottom). The level at which one signal neutrino

would be detected in total is indicated by the red dashed line. The  $S/\sqrt{S+B} = 1$  limit follows this line up until 10 detected BBH mergers. Afterwards, the background starts to become important. The improvement in  $f_{\text{BBH}}^{\nu}$  with localization by LIGO is indicated by the dashed and dashed-dotted blue lines. The expected number of BBH mergers in run O2 of LIGO is in the range 10-35 [24]. For the highest value of this range, it is possible to probe  $f_{\text{BBH}}^{\nu}$  down to about  $f_{\text{BBH}}^{\nu} \approx 8 \times 10^{-3}, 3 \times 10^{-3}, 5 \times 10^{-4}$  for  $S/\sqrt{S+B} = 5, 3, 1$  respectively. The astrophysical bounds corresponding to the merger rates in Eq. 2.4 are at  $f_{\text{BBH}}^{\nu} = 1.37 \times 10^{-3}$  and  $f_{\text{BBH}}^{\nu} = 5.15 \times 10^{-5}$ . The constraint on  $f_{\text{BBH}}^{\nu}$  reaches the weakest astrophysical bound at  $N_{\text{GW}} \gtrsim 180, 70, 8$  for  $S/\sqrt{S+B} = 5, 3, 1$ . A neutrino signal before 10 mergers would therefore violate the astrophysical bound, invalidating the assumptions in its calculation. This would therefore constrain the cosmic evolution of the BBH mergers or require the neutrino emission fraction  $f_{\text{BBH}}^{\nu}$  to be different for certain BBH merger classes. If a neutrino signal is only detected for higher  $N_{\text{GW}}$ , this will determine the level at which BBH mergers contribute to the diffuse astrophysical neutrino flux.



**Figure 2:** The expected limit on  $f_{\text{BBH}}^{\nu}$  for different signal strengths (from top to bottom  $S/\sqrt{S+B} = 5, 3, 1$ ) as a function of number of  $3 M_{\odot}$  BBH mergers observed in gravitational waves. The background is integrated over a sky patch of  $600 \text{ deg}^2$ ,  $100 \text{ deg}^2$  and  $20 \text{ deg}^2$  for the full, dashed and dashed-dotted lines respectively. The astrophysical bounds corresponding to the rates in Eq. 2.4 are shown by the green hatched lines. The band indicates the expected number of BBH mergers in LIGO run O2. Figure from [12].

## 5. Model-dependent interpretation

It is useful to interpret this general limit on the neutrino emission fraction  $f_{\text{BBH}}^{\nu}$  in more specific models of neutrino production. If it is assumed that no energy is extracted from the black holes or the gravitational waves themselves, but instead the neutrino emission originates from matter around the black hole binary, we can decompose  $f_{\text{BBH}}^{\nu}$  as

$$f_{\text{BBH}}^{\nu} = f_{\text{matter}} \times f_{\text{engine}} \times \epsilon_{\text{p,acc}} \times \epsilon_{\nu}. \quad (5.1)$$

The first part,  $f_{\text{matter}}$ , characterizes the amount of matter present in the black hole binary environment relative to the energy emitted in gravitational waves. The second part,  $f_{\text{engine}} \times \epsilon_{\text{p,acc}} \times \epsilon_{\nu}$ , characterises the acceleration of protons in order to create high energy neutrinos. The fraction of energy deposited from the surrounding matter in the acceleration engine is given by  $f_{\text{engine}}$ . The amount of energy going from the engine to the protons is given by  $\epsilon_{\text{p,acc}}$ , while the amount of energy going from protons into neutrinos is given by  $\epsilon_{\nu}$ .

If we assume that the acceleration of particles is similar to the case of a short GRB, we can use the GRB-fireball model [25] to calculate the acceleration part of Eq. 5.1. In this model, one typically finds that the accretion disk mass is converted into fireball energy with an efficiency  $f_{\text{engine}} = 1/10$ . The protons can receive a fraction  $\epsilon_{\text{p,acc}} = 1/10$  of the fireball energy and then produce neutrinos with  $\epsilon_{\nu} = 1/20$  of that energy. Therefore, we find that

$$f_{\text{BBH}}^{\nu} = f_{\text{matter}} \times 5 \cdot 10^{-4}. \quad (5.2)$$

This result can be used to put a bound on the amount of matter present in the binary black hole environment using the non-detection of counterpart neutrinos in GW150914. Since the bound in Section 3 was at  $f_{\text{BBH}}^{\nu} = 3.96 \times 10^{-5}$  for beamed emission in a typical solid angle  $\Delta\Omega = 0.2 \times 0.2$ , we estimate

$$f_{\text{matter}}^{\text{GW150914}} \lesssim 7.9 \times 10^{-2} \times \frac{\Delta\Omega}{0.2 \times 0.2}. \quad (5.3)$$

Using the expected limit at  $S/\sqrt{S+B} = 1$  on  $f_{\text{BBH}}^{\nu}$  after 35 BBH merger events detected by LIGO from the previous section, the expected limit on the amount of matter is

$$f_{\text{matter}}^{\text{N}_{\text{GW}}=35} \lesssim 3.2 \times 10^{-3} \times \frac{\Delta\Omega}{0.2 \times 0.2}. \quad (5.4)$$

If instead the bound on  $f_{\text{BBH}}^{\nu}$  from the diffuse astrophysical neutrino flux from mergers similar to GW150914 is used (Section 3), which is independent of beaming, we find

$$f_{\text{matter}}^{\text{astrophys}} \lesssim 2.3 \times 10^{-2}. \quad (5.5)$$

It is also possible to reverse the argument and find the neutrino emission fraction  $f_{\text{BBH}}^{\nu}$  for realistic models of the black hole binary environment. In case the emission is coming from a non-active accretion disk that is reactivated upon the merger [2], one expects about  $10^{-3}$ - $10^{-4} M_{\odot}$  of matter. The expected neutrino emission fraction is then

$$f_{\text{BBH}}^{\nu} \approx 10^{-7}, \quad (5.6)$$

which is still below the expected bound in the near future (Fig. 2).

## 6. Conclusion

Using the observation of the binary black hole merger GW150914 and its corresponding follow-up search for high energy neutrinos by IceCube, it is possible to put a limit on the neutrino emission from binary black hole mergers. We parametrise this emission with the neutrino emission fraction  $f_{\text{BBH}}^{\nu} = \frac{E_{\nu}}{E_{\text{GW}}}$ . Using the non-detection of counterpart neutrinos in this event, it is

possible to put a limit at  $f_{\text{BBH}}^{\text{V}} = 1.24 \times 10^{-2}$  for isotropic emission and  $f_{\text{BBH}}^{\text{V}} = 3.96 \times 10^{-5}$  for beamed emission in a patch  $\Delta\Omega = 0.2 \times 0.2$ . In addition, the observation of the diffuse astrophysical neutrino flux by IceCube limits the neutrino emission to  $f_{\text{BBH}}^{\text{V}} = 3.63 \times 10^{-3}$  (central value) from a class of mergers similar to GW150914.

The prospects for the expected limit on  $f_{\text{BBH}}^{\text{V}}$  were shown as a function of the number of future BBH mergers observed by LIGO. It is found that after about 10 mergers, the limit on  $f_{\text{BBH}}^{\text{V}}$  from direct searches will reach the limit set by the diffuse astrophysical neutrino flux at  $S/\sqrt{S+B} = 1$ . The expected number of events in LIGO run O2 is between 10-35. Using the higher value of this estimate, we find that the expected limit of  $f_{\text{BBH}}^{\text{V}}$  will reach  $f_{\text{BBH}}^{\text{V}} \approx 8 \times 10^{-3}, 3 \times 10^{-3}, 5 \times 10^{-4}$  for  $S/\sqrt{S+B} = 5, 3, 1$ .

Finally, we show how to interpret this generic bound on the neutrino emission fraction  $f_{\text{BBH}}^{\text{V}}$  in the context of specific models. Using the GRB-fireball model as the acceleration engine, we can constrain the amount of matter in the BBH merger environment to  $f_{\text{matter}}^{\text{GW150914}} \lesssim 7.9 \times 10^{-2} \times \frac{\Delta\Omega}{0.2 \times 0.2}$  from GW150914,  $f_{\text{matter}}^{\text{NGW}=35} \lesssim 3.2 \times 10^{-3} \times \frac{\Delta\Omega}{0.2 \times 0.2}$  from the expected limit on  $f_{\text{BBH}}^{\text{V}}$  after 35 events and  $f_{\text{matter}}^{\text{astrophys}} \lesssim 2.3 \times 10^{-2}$  from the astrophysical bound for mergers similar to GW150914. In order to give an estimate of  $f_{\text{BBH}}^{\text{V}}$  for realistic situations, we applied the same method to a model of a dead accretion disk around one of the black holes which is reactivated upon the merger. We then find  $f_{\text{BBH}}^{\text{V}} \approx 10^{-7}$ , which is still below the expected reach in LIGO run O2.

## Acknowledgments

KDV is supported by the Flemish Foundation for Scientific Research (FWO-12L3715N - K.D. de Vries). GDW and JMF are supported by Belgian Science Policy (IAP VII/37) and JMF is also supported in part by IISN. MV is aspirant FWO Vlaanderen.

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